

Complementarity and Searches for Dark Matter in the pMSSM

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Abstract

The search for and identification of neutralino dark matter in supersymmetry requires a multi-pronged approach with important roles played by collider, direct and indirect dark matter detection experiments. In this report, we summarize the sensitivity of such searches at the 7, 8 (and eventually 14) TeV LHC, combined with those by *Fermi*, CTA, IceCube/DeepCore, COUPP and XENON1T, to such particles within the context of the 19-parameter p(henomenological)MSSM. This report provides an outline of the current status of our results and our expectations for future analyses.

1 Introduction and Overview of the pMSSM

One of the main motivations for R-parity conserving supersymmetry (SUSY) is the prediction that the lightest SUSY particle (LSP) is stable and may be identified as a candidate thermal dark matter (DM) particle if it is both electrically neutral and colorless. Quite commonly the role of the LSP is played by the lightest neutralino, χ_1^0 , and this will be assumed in the discussion here. While DM searches are superficially focused on the nature of the LSP, the properties of all the other superparticles as well as those of the extended SUSY Higgs sector also come into play. Thus it is impossible to completely separate DM searches from searches for and the examination of the rest of the SUSY spectrum. However, even in the simplest SUSY scenario, the MSSM, the number of free parameters (~ 100) is too large to study in all generality. The traditional approach is to assume the existence of a high-scale

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theory with only a few parameters (such as mSUGRA [1]) from which all the properties of the sparticles at the TeV scale can be determined and studied in detail. While such an approach is quite valuable [2], these scenarios are somewhat phenomenologically limiting and are under increasing tension with a wide range of experimental data including, in some cases, the ~ 126 GeV mass of the recently discovered Higgs boson [3, 4].

One way to circumvent such limitations is to examine the far more general 19-parameter pMSSM [5]. The pMSSM is the most general version of the R-parity conserving MSSM when it is subjected to several experimentally-motivated constraints: (i) CP conservation, (ii) Minimal Flavor Violation at the electroweak scale, (iii) degenerate first and second generation sfermion masses, (iv) negligible Yukawa couplings and A-terms for the first two generations. In particular, no assumptions are made about physics at high scales, e.g., the nature of SUSY breaking, in order to capture electroweak scale phenomenology for which a UV-complete theory may not yet exist. Imposing the constraints (i)-(iv) decreases the number of free parameters in the MSSM at the TeV-scale from 105 to 19 for the case of a neutralino LSP, or 20 when the gravitino mass is included as an additional parameter when it plays the role of the LSP¹. We have recently begun a detailed study of the signatures for the pMSSM at the 7 and 8 (and eventually 14) TeV LHC, supplemented by input from DM experiments as well as from precision electroweak and flavor measurements [6, 7]. In this work, we study and compare the sensitivities of present and future experiments employing different approaches to searching for dark matter: collider production, direct and indirect detection. Independently of the LSP type we will not assume that the thermal relic density as calculated for the LSP necessarily saturates the WMAP/Planck value [8] to allow for the possibility of multi-component DM. For example, the axions introduced to solve the strong CP problem may make up a substantial amount of DM. The 19 pMSSM parameters and the ranges of values employed in our scans are listed in Table 1. The lower and upper limits employed in our scan ranges were chosen to be essentially consistent with Tevatron and LEP data and to have kinematically accessible sparticles at the LHC, respectively. To study the pMSSM, we generate many millions of model points in this space (using SOFTSUSY [9] and checking for consistency using SuSpect [10]), with each point corresponding to a specific set of values for these parameters. These individual models are then subjected to a large set of collider, flavor, precision measurement, dark matter and theoretical constraints [6]. Roughly 225k models with a neutralino LSP survive this initial selection and can then be used for further physics studies. (We have also obtained a model set of similar size for the case of gravitino LSPs.) Decay patterns of the SUSY partners and the extended Higgs sector are calculated using a modified version of SUSY-HIT [11].

In addition to these two large pMSSM model sets, we have recently generated a smaller, specialized neutralino LSP set of ~ 10 k ‘natural’ models all of which predict $m_h = 126 \pm 3$ GeV, have an LSP that *does* saturate the WMAP relic density and which produce values of fine-tuning (FT) better than 1% using the Barbieri-Giudice measure [12, 13]. This low-FT model set will also be used in the future as part of our completed study.

As a result of the upper limit employed in our scan ranges of 4 TeV, an upper limit

¹In this work we will limit our discussion to the case of neutralino LSPs

$m_{\tilde{L}(e)_{1,2,3}}$	100 GeV – 4 TeV
$m_{\tilde{Q}(q)_{1,2}}$	400 GeV – 4 TeV
$m_{\tilde{Q}(q)_3}$	200 GeV – 4 TeV
$ M_1 $	50 GeV – 4 TeV
$ M_2 $	100 GeV – 4 TeV
$ \mu $	100 GeV – 4 TeV
M_3	400 GeV – 4 TeV
$ A_{t,b,\tau} $	0 GeV – 4 TeV
M_A	100 GeV – 4 TeV
$\tan \beta$	1 - 60
$m_{3/2}$	1 eV–1 TeV (\tilde{G} LSP)

Table 1: Scan ranges for the 19 (20) parameters of the pMSSM with a neutralino (gravitino) LSP. The gravitino mass is scanned with a log prior. All other parameters are scanned with flat priors, though we expect this choice to have little qualitative impact on our results [14].

chosen to enable phenomenological studies at the 14 TeV LHC, the LSPs in our model sets are typically very close to being a pure electroweak eigenstate since the off-diagonal elements of the chargino and neutralino mass matrices are at most $\sim M_W$. Figure 1 shows some properties of the nearly pure neutralino LSPs (defined as a single electroweak eigenstate comprising over 90% of the mass eigenstate). In the left panel we see the distribution of the LSP mass for nearly pure bino, wino, and Higgsino LSPs while in the right-hand panel we see the corresponding distribution for the predicted LSP thermal relic density. Note that the masses of all of our neutralino LSPs lie below ~ 2 TeV due to our choice of the scan ranges, and the entire SUSY spectrum must be heavier than the LSP and less than ~ 4 TeV by definition. The fraction of models which are nearly pure bino is found to be rather low in this model set since pure binos generally predict too high a relic density unless they co-annihilate with another sparticle, happen to be close to some (Z, h, A) funnel region or have a suitable Higgsino admixture. Note that only in the rightmost bin of the right panel is the relic density approximately saturating the WMAP thermal relic value. These LSP properties will be of particular importance in the discussion that follows.

Figure 2 shows the thermal relic densities of the LSPs in our pMSSM model sample as a function of their mass with the color-coding reflecting their electroweak eigenstate content. There are many things to note here that will be important for later consideration; essentially every possible mechanism to obtain (or lie below) the WMAP relic density is present: *(i)* The set of models at low masses on the left-hand side forming ‘columns’ correspond to bino and bino-Higgsino admixtures surviving due to their proximity to the Z, h -funnels. *(ii)* The bino-Higgsino LSPs saturating the relic density in the upper left are of the so-called ‘well-tempered’ variety. *(iii)* the pure² bino models in the middle top of the Figure are bino

²Here again, ‘pure’ means having an eigenstate fraction $\geq 90\%$. Points shown as bino-wino, bino-Higgsino, or wino-Higgsino mixtures have less than 2% Higgsino, wino, or bino fraction, respectively. Mixed points have no more than 90% and no less than 2% of each component.

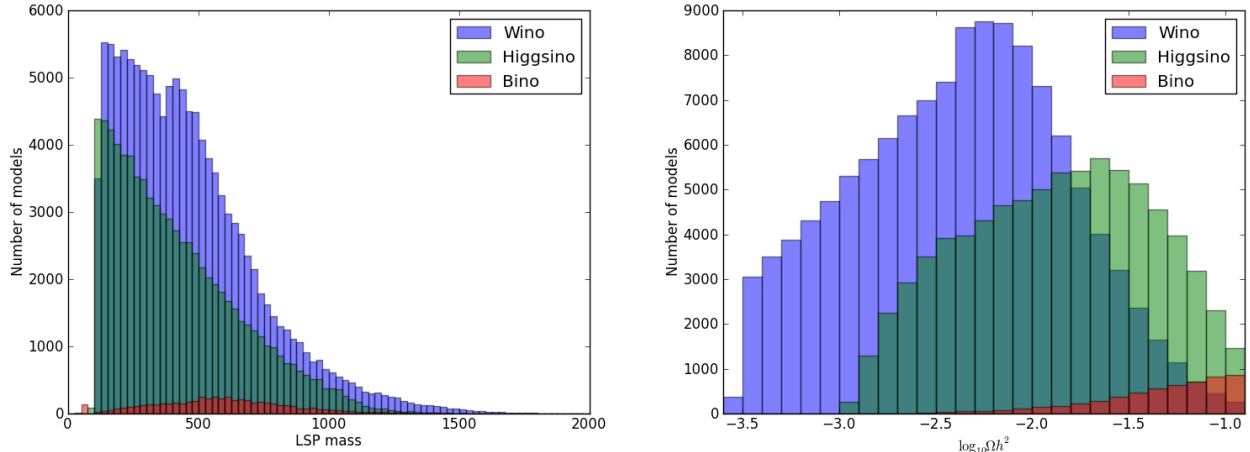


Figure 1: Distribution of the LSP masses (left) and predicted relic densities (right) for neutralino LSPs which are almost pure weak eigenstates.

co-annihilators (mostly with sleptons) or are models near the A -funnel region. (iv) The green (blue) bands are pure Higgsino (wino) models that saturate the relic density bound near $\sim 1(1.7)$ TeV but dominantly appear at far lower relic densities. Wino-Higgsino hybrids are seen to lie between these two cases as expected. (v) A smattering of other annihilator models are seen to be loosely distributed in the lower right-hand corner of the figure.

2 LHC Searches

We begin with a short overview of the searches for the pMSSM at the 7 and 8 TeV LHC [6, 7]. In general, we follow the suite of ATLAS SUSY analyses but also supplement these with several searches performed by CMS; the analyses are briefly summarized in Table 2. The SUSY searches are then further augmented by additional searches for heavy neutral SUSY Higgs $\rightarrow \tau^+\tau^-$ by CMS [15] and for the rare decay mode $B_s \rightarrow \mu^+\mu^-$ as discovered by LHCb [16], both of which play distinct but important roles in restricting the pMSSM parameter space. Presently, we have implemented every relevant ATLAS SUSY search available as of mid-September 2012. This is currently being updated to include all searches made public by ATLAS before the beginning of March 2013; even more recently released ATLAS analyses will be implemented by us during the late spring and summer of 2013 or as soon as they are made public.

Briefly stated our procedure is as follows: We generate SUSY events for each model using PYTHIA 6.4.26 [17] and PGS 4 [18], which we have modified to, e.g., correctly deal with gravitinos, hadronization of stable colored sparticles, multi-body decays and ATLAS b-tagging. We then scale our event rates to NLO by calculating the relevant K-factors using Prospino 2.1 [19]. The individual searches are then implemented using a custom analysis code [14], following the published cuts and selection criteria as closely as possible. This

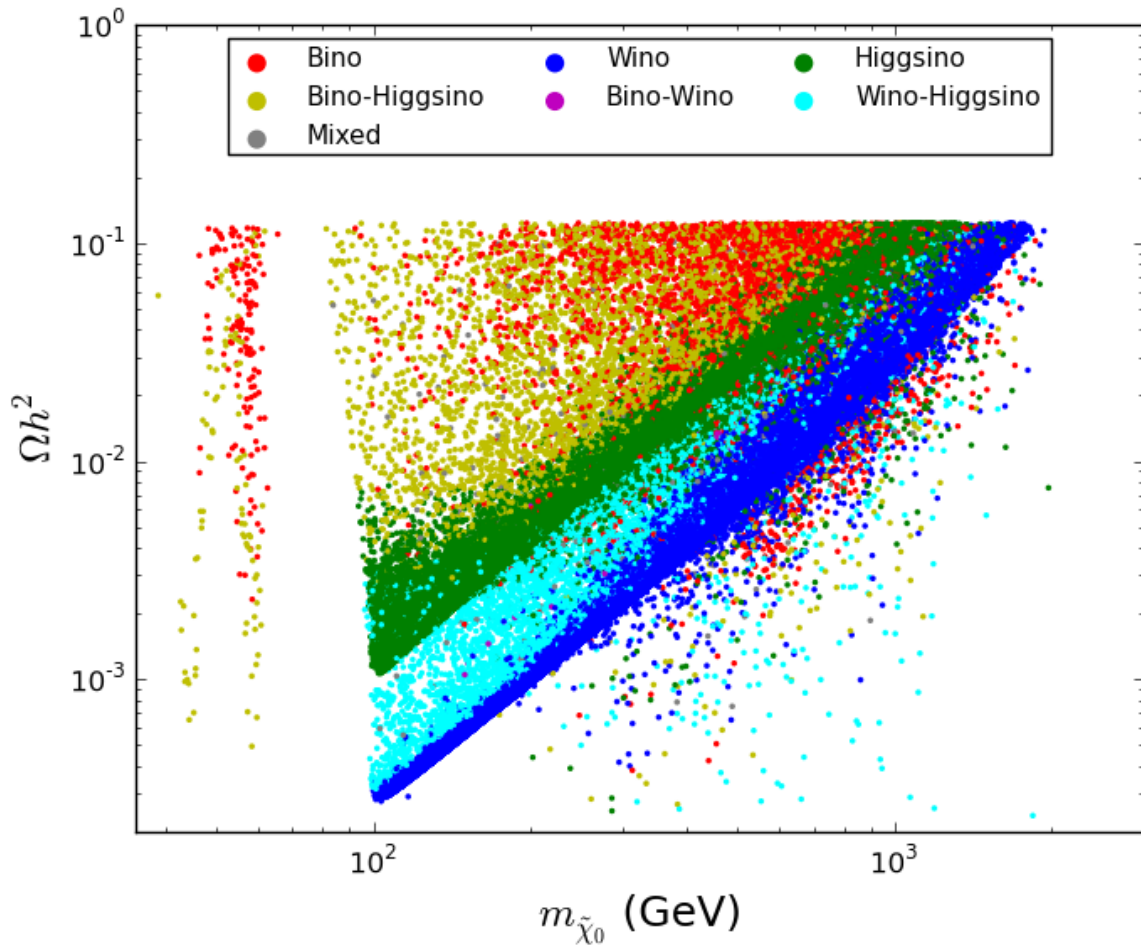


Figure 2: Thermal relic density as a function of the LSP mass in our pMSSM model set, as generated, color-coded by the electroweak properties of the LSP as discussed in the text.

Search	Energy	Reference
2-6 jets	7 TeV	ATLAS-CONF-2012-033
multijets	7 TeV	ATLAS-CONF-2012-037
1 lepton	7 TeV	ATLAS-CONF-2012-041
2-6 jets	8 TeV	ATLAS-CONF-2012-109
multijets	8 TeV	ATLAS-CONF-2012-103
1 lepton	8 TeV	ATLAS-CONF-2012-104
SS dileptons	8 TeV	ATLAS-CONF-2012-105
Glauino \rightarrow stop/sbottom	7 TeV	1207.4686
Very light stop	7 TeV	ATLAS-CONF-2012-059
Medium stop	7 TeV	ATLAS-CONF-2012-071
Heavy stop (0l)	7 TeV	1208.1447
Heavy stop (1l)	7 TeV	1208.2590
GMSB direct stop	7 TeV	1204.6736
Direct sbottom	7 TeV	ATLAS-CONF-2012-106
3 leptons	7 TeV	ATLAS-CONF-2012-108
1-2 leptons	7 TeV	1208.4688
Direct slepton/gaugino (2l)	7 TeV	1208.2884
Direct gaugino (3l)	7 TeV	1208.3144
HSCP	7 TeV	1205.0272
Disappearing tracks	7 TeV	ATLAS-CONF-2012-111
$\gamma\gamma + MET$ (\tilde{G} LSP)	7 TeV	1209.0753

Table 2: Simulated LHC SUSY searches which have been applied to our pMSSM model sets. 66 (53) percent of our models with a neutralino (gravitino) LSP are not excluded by any of these searches and this is found to be approximately independent of the Higgs mass constraint.

analysis code is validated for each of the many search regions in every analysis employing the benchmark model points provided by ATLAS (and CMS). Models are then excluded using the same CL_s criterion as is used by ATLAS. These analyses are performed *without* imposing the Higgs mass constraint, $m_h = 126 \pm 3$ GeV (combined experimental and theoretical errors) so that we can understand its influence on the search results. Note that roughly 20% of the neutralino model set (before the LHC SUSY searches are applied) predict a Higgs mass in the above range. While there is some variation amongst the individual searches we find that, once combined, the total fraction of our models surviving (or killed by) the set of all LHC searches is to an excellent approximation *independent* of whether or not the Higgs mass constraint has been applied. Conversely, the $\sim 20\%$ fraction of neutralino models predicting the correct Higgs mass is also found to be approximately independent of whether or not the SUSY searches have been applied. This combined result is very powerful, demonstrating the current approximate decoupling of SUSY search results from the discovery of the Higgs boson and allowing us to continue examining the entire model set with some validity.

Figure 3 shows some results demonstrating the power of this combination of LHC analyses to constrain the set of pMSSM neutralino LSP models. In particular this figure shows the fraction of models with a given sparticle and LSP mass which are excluded by the combined LHC searches. Since the masses of the squarks and gluinos, the lightest stops and sbottoms and the LSP itself are currently of particular interest, we have concentrated on those quantities in this figure. In the upper left-hand panel the coverage of the gluino-LSP mass plane by the LHC searches is displayed; the solid white line represents the Simplified Model 95% CL search limit from ATLAS [20] which we see is roughly the same as the all-black region excluded in the pMSSM. Interestingly a respectable number of pMSSM models are seen to survive below this boundary. The upper right-hand panel shows the corresponding coverage in the gluino-lightest squark mass plane with the analogous Simplified Model line from ATLAS [20] (which assumes a zero mass LSP and that the 8 squarks of the first two generations are degenerate, neither of which are necessarily true in the pMSSM). While most models with rather light squarks and/or gluinos are observed to be excluded by the LHC searches, it is clear that models with both squarks and gluinos below ~ 700 GeV still remain viable. The lower two panels show the lightest stop/sbottom vs. LSP mass plane with the corresponding ATLAS Simplified Model search limits [21]. Here we see several things, the most important being that the amount of coverage of these two mass planes in the pMSSM by the LHC searches is significantly larger than what we would have expected from either Simplified Model analysis. We also find that the third generation searches are ineffective for models that have a relatively small splitting between the stop and LSP ($\delta m \lesssim 130$ GeV), so the exclusion of these models relies almost entirely on the standard jets + MET searches. In the uncompressed regions, on the other hand, the third generation searches significantly outperform jets + MET, as expected. In particular, the $2b$ +MET analysis is found to be responsible for nearly all of the exclusion power in this region. Clearly, the third-generation searches cover important gaps in the standard jets + MET coverage. In general, the third-generation searches are most sensitive to models in which stop or sbottom decays produce hard b -jets, and have a more difficult time observing models with stops or sbottoms that produce tops or decay to multiple intermediate gauginos (resulting in softer b -jets and leptons). Lastly we again stress that the direct sbottom search is found to be responsible for the majority of the heavy flavor search exclusion power for models with relatively light stops.

In addition to the 7 and 8 TeV LHC searches, future data taking and enhanced analyses at ~ 14 TeV will greatly extend the expected coverage of the pMSSM parameter space. We hope to provide some estimate of this coverage later this summer using input from both the ATLAS and CMS collaborations. As of now we can only make a very rough guess at this coverage by making use of some scaling arguments, extrapolating our observed reach to higher energy and luminosity using power laws that are derived from ATLAS estimates. These results are shown in Fig. 4 for both 0.3 and 3 ab^{-1} of integrated luminosity in the gluino-lightest squark mass plane. Here we see that, for the equal mass case, one might expect that $\sim 70\%$ of the models would be covered up to squark/gluino masses of $\sim 2(2.5)$ TeV for these values of the integrated luminosity. As always, the most interesting models

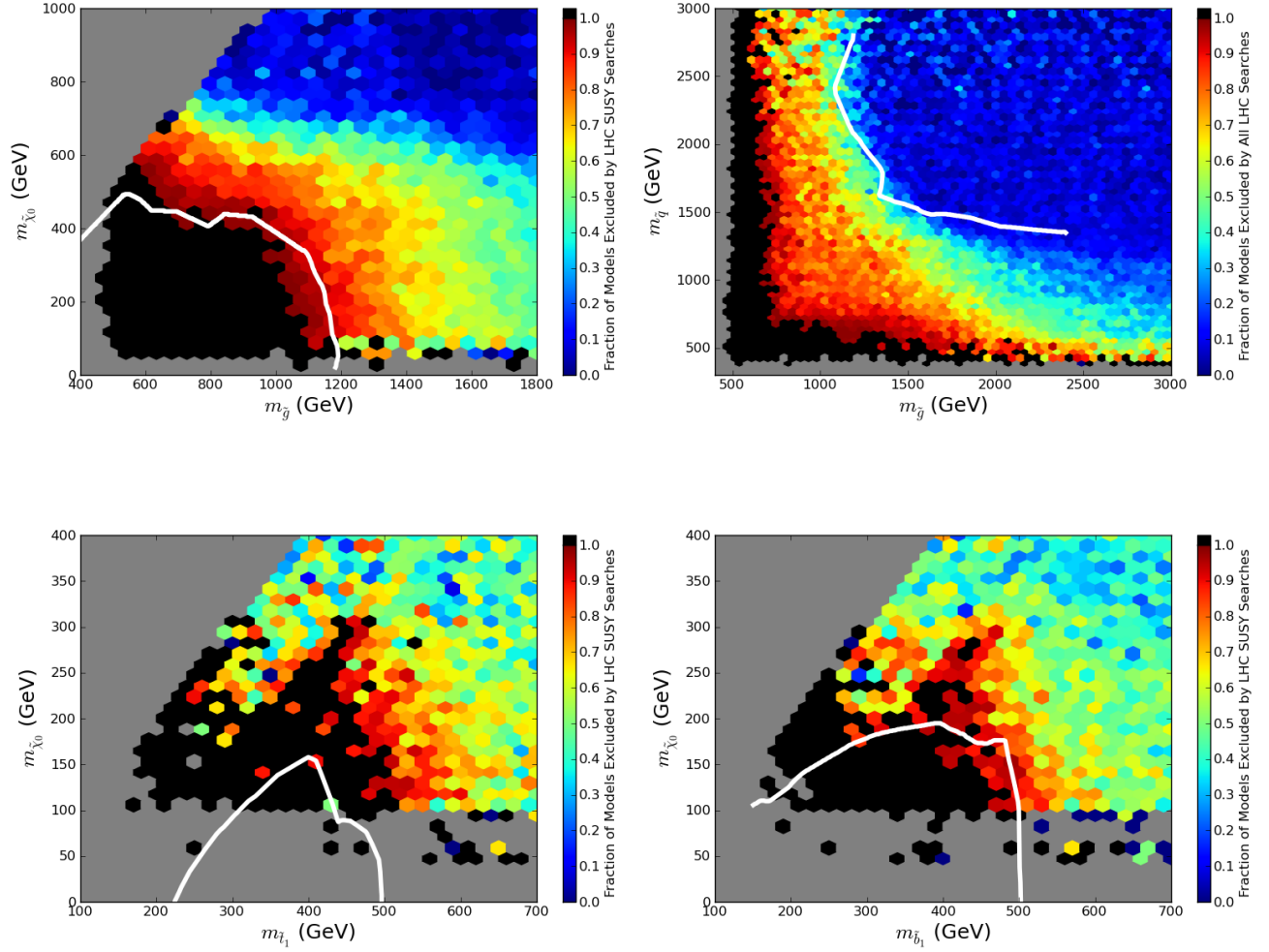


Figure 3: Projections of the pMSSM model coverage efficiencies from the 7 and 8 TeV LHC searches shown in the gluino-LSP mass plane (top left), the gluino-lightest squark mass plane (top right), the lightest stop-LSP mass plane (lower left) and in the lightest sbottom-LSP mass plane. The solid lines represent the corresponding 95% CL limit results obtained by ATLAS in the simplified model limit as discussed in the text.

will be the ones with relatively light sparticles which are nevertheless missed by the full set of analyses. To fully ascertain the nature of such models will require a far more realistic study than is presently available since the LHC collaborations themselves are still unsure of how their detectors will perform at 14 TeV with very high pileup conditions.

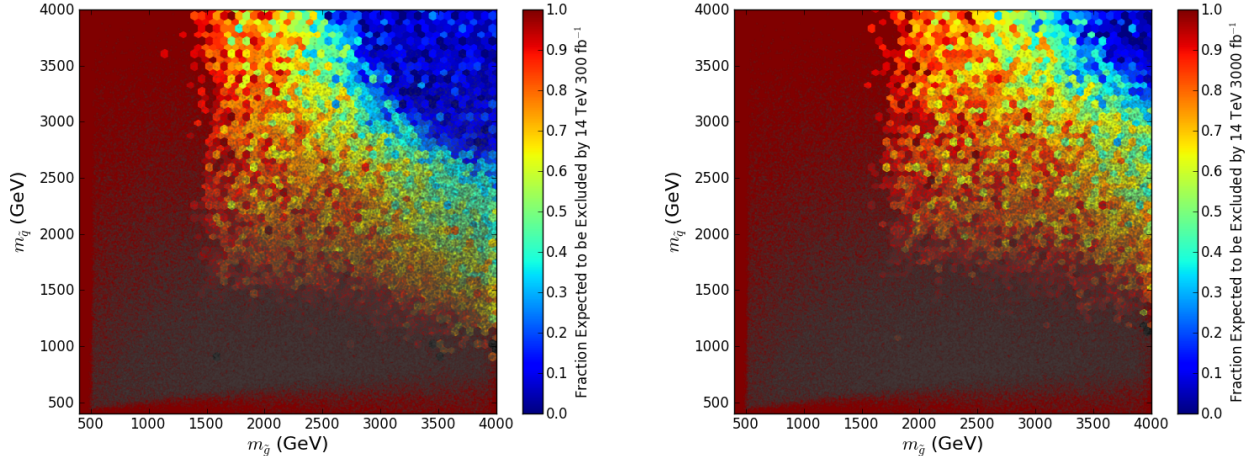


Figure 4: Guesstimates of the pMSSM model coverage efficiency at the 14 TeV LHC projected onto the gluino-lightest squark plane assuming 0.3 (left) or 3 ab^{-1} (right) of integrated luminosity. To obtain these estimates, we have scaled our observed 8 TeV search efficiency results to higher energy and luminosity. Our models are overlaid as gray points.

3 Direct Detection

The direct detection of DM results from either the spin-independent (SI) or spin-dependent (SD) scattering of the LSP by a target nucleus. While $Z(h)$ t -channel exchange only contributes to the SD(SI) process at tree level, s -channel squark exchange can contribute to both processes. However, as the lower bounds on the first and second generation squark masses from the LHC become stronger the importance of these squark exchange contributions clearly become subdominant.³ The Z -exchange graph is sensitive to the Higgsino content of the LSP whereas the Higgs exchange graph probes the product of the LSP's gaugino and Higgsino content. Similarly, (valence) squark exchange is sensitive to the LSP's wino and bino content.

Figure 5 displays the predicted experimentally observable SI and SD cross sections for our pMSSM model set together with several present and anticipated future experimental constraints. Note that the actual cross sections are appropriately scaled by the factor $R =$

³It is important to note, however, that since the many different squark masses can vary independently in the pMSSM and that the LHC constraints on, e.g., the u_L, d_L, u_R and d_R squarks are quite different, both SD and SI interactions may have substantial isospin-dependent contributions so that σ_p and σ_n can be significantly different.

$\Omega h^2/\Omega h^2|_{WMAP}$ to account for the fact that most of our pMSSM models lead to a thermal relic density somewhat below that which saturates the WMAP value as discussed in the Introduction. There are several things to note in this Figure: (i) Future SI searches will cut rather deeply into the model set; a lack of signal at XENON1T (LZ) would exclude 23% (50%) of these models. *However*, this implies that about half of our models are not accessible to SI experiments due to their rather small scaled cross sections, $R\sigma$. Models with these values of $R\sigma$ tend to produce these small values due to both the suppression arising from their low thermal relic density as well as fact that the LSPs tend to be rather close to being pure weak eigenstates as discussed above. Note that if we only consider models which predict a relic density within 10% of the critical density, the situation improves significantly, with 60% (87%) of models within the XENON1T (LZ) exclusion limit. (ii) SD measurements are still rather far from the pMSSM model predictions and we do not expect future SD measurements to have a significant impact on the parameter space; SD experiments such as COUPP500 will only be able to exclude $\sim 2\%$ of the models in this set if no signal is seen.

While direct detection experiments have significant power to cover much of the pMSSM, clearly such experiments will need to be ‘supplemented’ if we want to discover or exclude the full range of neutralino LSPs in the pMSSM model space. In the next two sections, we describe constraints on the pMSSM from indirect detection and neutrino telescope experiments.

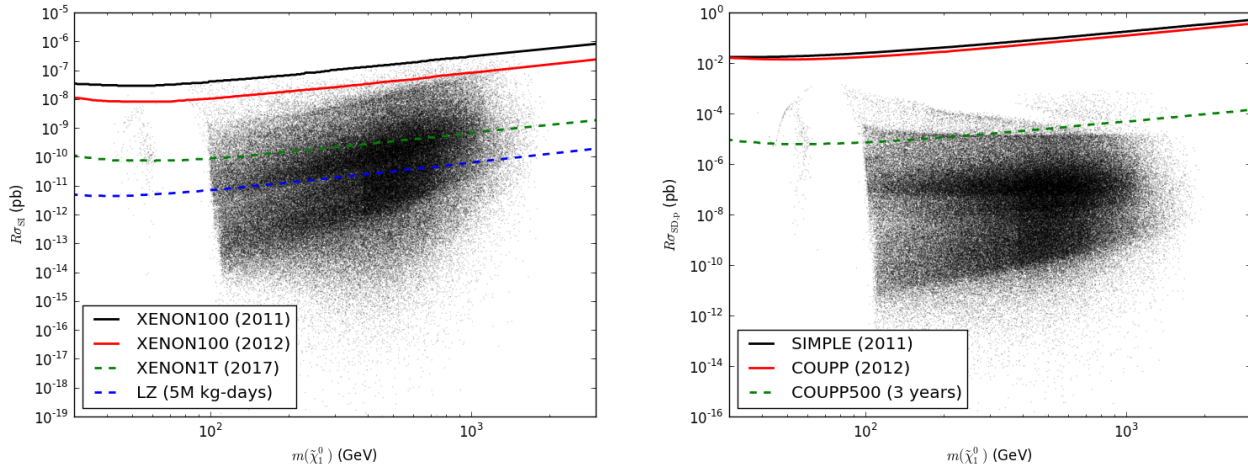


Figure 5: Scaled spin-independent (left) and spin-dependent (right) direct detection cross sections for our neutralino LSPs in comparison to current and future experimental sensitivities. The scaling factor accounts for the possibility that the calculated thermal relic density of the LSP is below that measured by WMAP.

4 Indirect Detection: *Fermi* LAT and CTA

Indirect Detection plays a critical role in searches for DM and, in the case of null results, can lead to very strong constraints on the pMSSM parameter space. As will be seen below, both *Fermi* and CTA can contribute in different regions of the pMSSM parameter space in the future. CTA, in particular, will be seen to be extremely powerful in the search for heavy LSPs which are mostly Higgsino- or wino-like and that predict thermal relic densities within an order of magnitude of the WMAP/Planck value. *Fermi*, on the other hand, will be seen to be mostly sensitive to well-tempered neutralinos that are relatively light.

The most promising DM targets for both *Fermi* LAT and CTA are those with both a high DM density and low astrophysical gamma-ray foregrounds. These criteria have motivated a number of Galactic and extragalactic targets including the Galactic Center (GC), dwarf spheroidal galaxies (dSphs), and galaxy clusters. The expected gamma-ray signal for DM annihilations is proportional to the integral of the square of the DM density along the line of sight to the source (J). The determination of J is most reliable for DM-dominated objects such as dSphs and galaxy clusters in which the DM distribution can be robustly measured. In the Milky Way halo the uncertainty on the DM distribution rapidly increases as one approaches the inner galaxy where baryons dominate the gravitational potential. Kinematic data at large scales constrain the density of DM at the solar radius to 0.2–0.4 GeV cm⁻³ [22, 23].

Under well-motivated models for the DM distribution in the Galactic halo, the GC is expected to be the most intense DM source in the sky. CDM simulations predict that the Galactic DM halo should have a density profile with an inner cusp, $\rho \propto r^{-\gamma}$ with $\gamma \simeq 1.0$. For an extrapolation of the Galactic DM density profile with $\rho(r) = \rho_{\odot}(r/r_{\odot})^{-1}$, the expected DM signal from the GC region is approximately two orders of magnitude greater than dSphs or galaxy clusters. However the GC is also the region of the sky with the highest density of gamma-ray sources and brightest diffuse gamma-ray emission produced from the interaction of cosmic rays with the interstellar medium. These foregrounds significantly limit the sensitivity of the *Fermi*-LAT in the inner galaxy and complicate the interpretation of any observed signals. For our study of *Fermi*-LAT we focus on the sensitivity to the DM signal from dSphs which are predominantly at high galactic latitudes where the astrophysical foregrounds are much weaker. Above 50 GeV the diffuse emission from the Galaxy is much less intense relative to other backgrounds than in the *Fermi*-LAT sensitivity band. The inner galaxy is thus the preferred target for CTA under the assumption that the Galactic DM halo possesses an inner cusp with $\gamma \simeq 1.0$.

4.1 *Fermi* LAT

Here we follow the procedure developed in Ackermann et al. (henceforth A11) [24] and extended in Cotta et al. [25] to constrain the annihilation cross section, $\langle\sigma v\rangle$, for each pMSSM model using *Fermi*-LAT observations of ten dwarf spheroidal galaxies (dSphs). Our two-year γ -ray event sample is identical to that described in A11, accepting photons in the energy range from 200 MeV $< E <$ 100 GeV within 10° of each dSph. In accord with

A11, we use the LAT ScienceTools⁴ version v9r20p0 and the P6.V3.DIFFUSE instrument response functions.⁵ J-factors and associated statistical uncertainties for the dSphs are taken from Table 1 of A11, where they were calculated using line-of-sight stellar velocities and the Jeans equation [24]. Similar to Cotta et al., we use DarkSUSY 5.0.5 [26] to model the γ -ray spectrum from the annihilation of each pMSSM LSP. DarkSUSY calculates the total γ -ray yield from annihilation, as well as the rates into each of 27 final state channels.

We calculate a joint likelihood to constrain the annihilation cross section of each pMSSM model given the LAT observations coincident with the ten dSphs. Following A11, we incorporated statistical uncertainties in the J-factors of the dSphs as nuisance parameters in our likelihood formulation (see Equation 1 and the associated discussion in A11). This likelihood formulation includes both the flux normalizations of background γ -ray sources (diffuse and point-like) and the associated dSph J-factors and statistical uncertainties. No significant γ -ray signal is detected from any of the dSphs when analyzed individually or jointly for any of the pMSSM models.

For each of the pMSSM models, we calculate the maximum annihilation cross section, $\langle\sigma v\rangle_{\text{UL}}$, consistent with the null detection in the LAT data. We incorporate nuisance parameters to obtain a 95% one-sided confidence interval on the value of $\langle\sigma v\rangle$ using the profile likelihood method [27]. This one-sided 95% confidence limit on $\langle\sigma v\rangle$ serves as our value of $\langle\sigma v\rangle_{\text{UL}}$, which is compared to the true value of the annihilation cross section for each pMSSM model. We define the “boost” necessary to constrain a model as the ratio $\langle\sigma v\rangle_{\text{UL}}/\langle\sigma v\rangle$.

While the LAT data do not presently constrain any of the pMSSM models, it is useful to estimate the improvements expected over a 10 year mission lifetime. In the low-energy, background dominated regime, the LAT point source sensitivity increases as roughly the square-root of the integration time. However, in the high-energy, limited background regime (where many pMSSM models contribute), the LAT sensitivity increases more linearly with integration time. Thus, 10 years of data could provide a factor of $\sqrt{5}$ to 5 increase in sensitivity. Additionally, optical surveys such as Pan-STARRS [28], the Dark Energy Survey [29], and LSST [30] could provide a factor of 3 increase in the number of Milky Way dSphs corresponding to an increased constraining power of $\sqrt{3}$ to 3 [31]. Ongoing improvements in LAT event reconstruction, a better understanding of background contamination, and an increased energy range are all expected to provide additional increases in the LAT sensitivity. Thus, we find it plausible that the LAT constraints could improve by a factor of 10 compared to current constraints.

In Figure 6 we display the boost required to constrain the various pMSSM models at 95% CL based on the *Fermi*-LAT dwarf analysis employing only the first 2 years of data colored-coded by either the annihilation cross section or the LSP thermal relic density. Here we see that the LAT analysis does not currently constrain any of our pMSSM models. However, as discussed above with more dwarfs and longer integration times we would expect an ~ 10 -fold improvement in the sensitivity and thus all models with boost factors less than 10 would become accessible. We will assume this ~ 10 -fold improvement in sensitivity assumption for

⁴<http://fermi.gsfc.nasa.gov/ssc/data/analysis/software>

⁵<http://fermi.gsfc.nasa.gov/ssc/data/analysis/scitools/overview.html>

the analysis that follows.

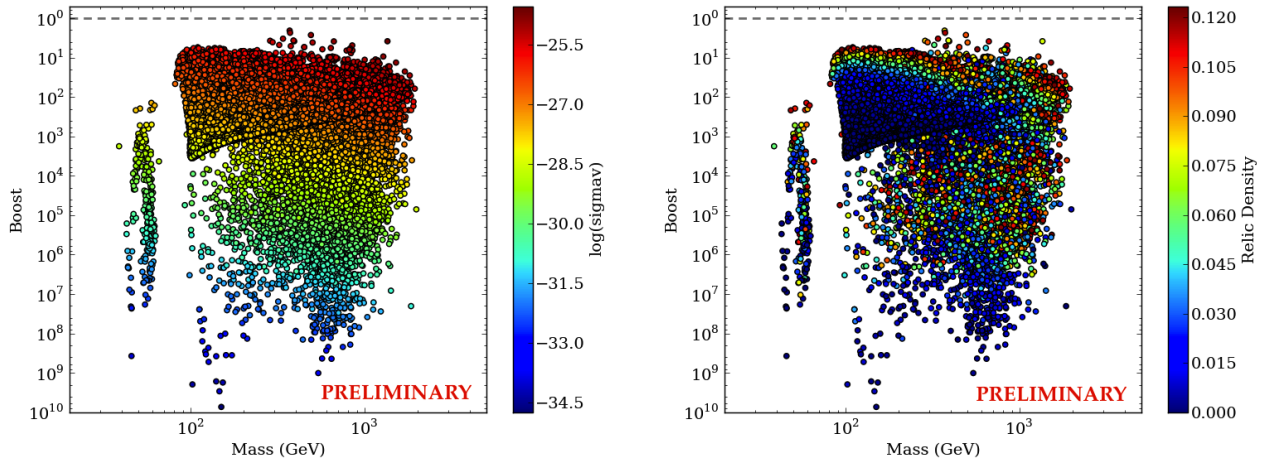


Figure 6: (Left) *Fermi*-LAT boost factor vs. LSP mass for the pMSSM model set. The true cross section, $\langle\sigma v\rangle$, for each model is plotted on the color scale. (Right) Here the corresponding relic density for each model is plotted on the color scale.

4.2 CTA

The Cherenkov Telescope Array (CTA) [32] is a future ground-based gamma-ray observatory that will have sensitivity over the energy range from a few tens of GeV to a few hundreds of TeV. To achieve the best sensitivity over this wide energy range CTA will include three telescope types: Large Size Telescope (LST, 23 m diameter), Medium Size Telescope (MST, 10-12 m) and Small Size Telescope (SST, 4-6 m). Over this energy range the point-source sensitivity of CTA will be at least one order of magnitude better than current generation imaging atmospheric Cherenkov telescopes such as H.E.S.S., MAGIC, and VERITAS. CTA will also have an angular resolution at least 2–3 times better than current ground-based instruments, improving with energy from 0.1° at 100 GeV to better than 0.03° at energies above 1 TeV.

The optimal DM search region for CTA will be limited by the CTA FoV of 8° to the area within 2° to 3° of the GC. On these angular scales the DM signal is predominantly determined by the DM distribution in the inner galaxy ($R_{GC} < 1$ kpc). We model the Galactic DM distribution with an NFW profile with a scale radius of 20 kpc normalized to 0.4 GeV cm^{-3} at the solar radius. This model is consistent with all current observational constraints on the Galactic DM halo and represents a conservative expectation for the inner DM profile in the absence of baryons. Because the annihilation signal is proportional to the square of the DM density, the projected limits for CTA depend strongly on the assumptions that are made on the shape and normalization of the Galactic DM halo profile. The projected limits presented here could change by as much as a factor of 10 given these uncertainties.

The potential of CTA for DM searches and testing other exotic physics has been studied in detail by [33] using the projected performance of several alternative array configurations [34] with 18–37 MSTs and different combinations of SSTs and LSTs. For this study we use the projected sensitivity of a candidate CTA configuration with 61 MSTs on a regular grid with 120 m spacing [35]. This telescope layout is similar to the anticipated layout of the baseline MST array with 18–25 telescopes with an additional US contribution of 36 MSTs. This configuration has a gamma-ray angular resolution that can be parameterized as a function of energy as $\theta \simeq 0.07^\circ (E/100 \text{ GeV})^{-0.5}$ and a total effective area above 100 GeV of $\sim 10^6 \text{ m}^2$. We define the GC signal region as an annulus centered on the GC that extends from 0.3° – 1.0° and calculate the sensitivity of CTA for an integrated exposure of 500 hours that is uniform over the whole region. An energy-dependent model for the background in the signal region is taken from a simulation of residual hadronic contamination. The uncertainty in the background model is calculated for a control region with no signal contamination and a solid angle equal to five times the signal region (14.3 deg^2).

We compute the sensitivity of CTA using a two-dimensional likelihood analysis that models the distribution of detected gamma-rays in energy and angular offset from the GC. Following the same procedure as the *Fermi*-LAT analysis, we compute the model boost factor as the ratio of the model cross section with the 95% C.L. upper limit on the annihilation cross section of a DM model with the same spectral shape. For models with LSP masses below 1 TeV, the sensitivity of CTA is dominated by events at low-energy ($E \lesssim 300 \text{ GeV}$).

Figure 7 shows the distribution of the CTA boost factor versus LSP mass for all pMSSM models and the subset of models that have an LSP relic density consistent with 100% of the DM relic density. CTA can exclude $\sim 20\%$ of the total model set and $>50\%$ of the models in the subset with the DM relic density. Here we see that $\sim 19\%$ of the models would be excluded by CTA if no signal were to be observed.

5 IceCube

Neutralino dark matter can be captured and accumulated in the sun. Neutralinos in this relatively overdense population would then also sink to the solar core and annihilate. If the product of capture and annihilation cross sections is large enough this process leads to an equilibrium population of captured neutralinos whose annihilations are proportional to their *elastic scattering cross-sections* [36, 37] and that may be detectable by observing an excess of high-energy ($\geq \text{GeV}$) solar neutrinos in km^3 -scale neutrino telescope experiments.

Here we present predictions for an IceCube/DeepCore (IC/DC) search for neutralino DM in our model set. Our analysis closely follows that presented in [38]. In the results presented here we assume that each neutralino’s relic density is given by the usual thermal calculation. We use DarkSUSY 5.0.6 [26] to simulate the (yearly-average) signal $\nu_\mu/\bar{\nu}_\mu$ neutrino flux spectra incident at the detector’s position and convolve with preliminary $\nu_\mu/\bar{\nu}_\mu$ effective areas for muon events contained in DeepCore⁶. We consider a data set that includes $\sim 5\text{yr}$

⁶These are the same effective areas that were used in [38], referred to there as “SMT8/SMT4.”

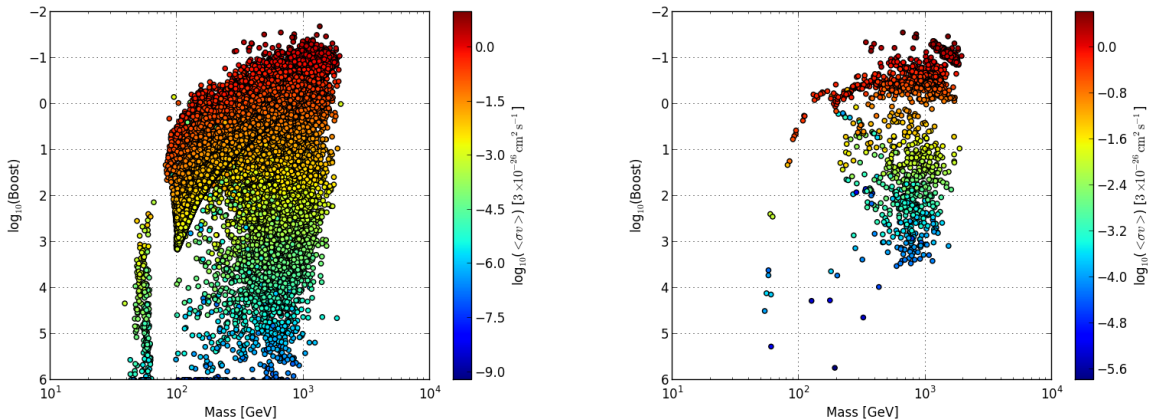


Figure 7: CTA boost factor vs. LSP mass for the full pMSSM model set (left) and pMSSM models satisfying current experimental constraints and a LSP relic density consistent with the WMAP7 DM density. The true cross section, $\langle\sigma v\rangle$, for each model is plotted on the color scale.

of data that is taken during austral winters (the part of the year for which the sun is in the northern hemisphere) over a total period of ~ 10 yrs ⁷. An irreducible background rate of ~ 10 events/yr is expected from cosmic ray interactions with nuclei in the sun. Here we will take (as discussed at greater length in [38]) a detected flux of $\Phi = 40$ events/yr as a conservative criterion for exclusion.

The basic results of this analysis are presented in Figure 8. In Figure 8, we show all pMSSM models in our set using grey points, while highlighting WMAP-saturating models with mostly bino, wino, Higgsino or mixed ($\leq 80\%$ of each) LSPs in red, blue, green and magenta, respectively. Detectability is tightly correlated with the elastic scattering cross-sections (σ_{SI} and σ_{SD}) while having little correlation with the annihilation cross-section $\langle\sigma v\rangle$, as expected.

The biggest difference between these results and those of the previous analysis [38], which used a set of pMSSM models that were chosen to have relatively light (≤ 1 TeV) sparticles, is that a much smaller percentage of the current pMSSM models are able to reach capture/annihilation equilibrium in the sun. This is due to the fact that so many of these models are nearly pure wino or Higgsino gauge eigenstates (which have both low relic density and small capture cross-sections) and that the LSPs in this model set tend to be much heavier than those in the previous set. If one defines out-of-equilibrium models as those with solar annihilation rates less than 90% of their capture rates, we find that no such models can be excluded by IC/DC. In contrast, relatively light LSPs composed of a mixture of gaugino and Higgsino eigenstates have large scattering and annihilation cross sections and

⁷In practice the IC/DC treatment of data is more sophisticated, classifying events as through-going, contained and strongly-contained, and allowing for some contribution from data taken in the austral summer. We expect that inclusion of this data would affect our results at a quantitative, but not qualitative, level.

are highly detectable by IC/DC. We observe that all such WMAP-saturating well-tempered neutralinos with masses $m_{LSP} \leq 500$ GeV should be excluded by the IC/DC search (*c.f.*, the magenta points in Fig. 8).

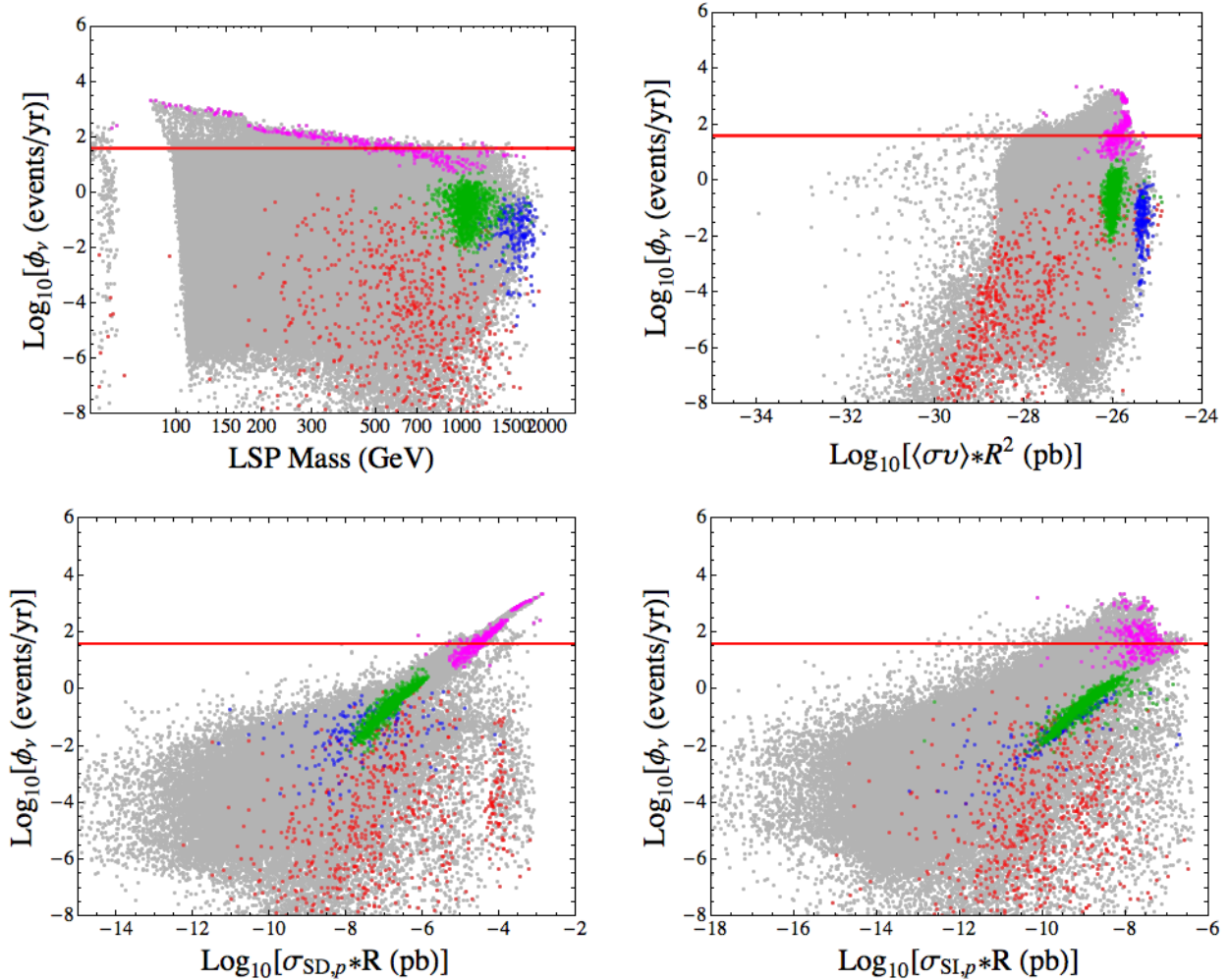


Figure 8: IC/DC signal event rates as a function of LSP mass (upper-left), thermal annihilation cross-section $\langle\sigma\nu\rangle R^2$ (upper-right) and thermal elastic scattering cross-sections $\sigma_{SD,p}$ and $\sigma_{SI,p}$ (lower panels). In all panels the gray points represent generic models in our full pMSSM model set, while WMAP-saturating models with mostly bino, wino, Higgsino or mixed ($\leq 80\%$ of each) LSPs in are highlighted in red, blue, green and magenta, respectively. The red line denotes a detected flux of 40 events/yr, our conservative estimate for exclusion.

6 Complementarity: Putting It All Together

Now that we have provided an overview of the various pieces of data that go into our analysis, we can put them together to see what they (will) tell us about the nature of the neutralino

LSP as DM and, more generally, the pMSSM itself. We remind the reader that this is an ongoing analysis and that several future updates will be made to what we present here before completion. In particular, the LHC analyses will require updating to include more results at 8 TeV along with our extrapolations to 14 TeV. While these are important pieces to the DM puzzle it is our expectation that the addition of these new LHC results will only strengthen the important conclusions based on the existing analyses to be discussed below.

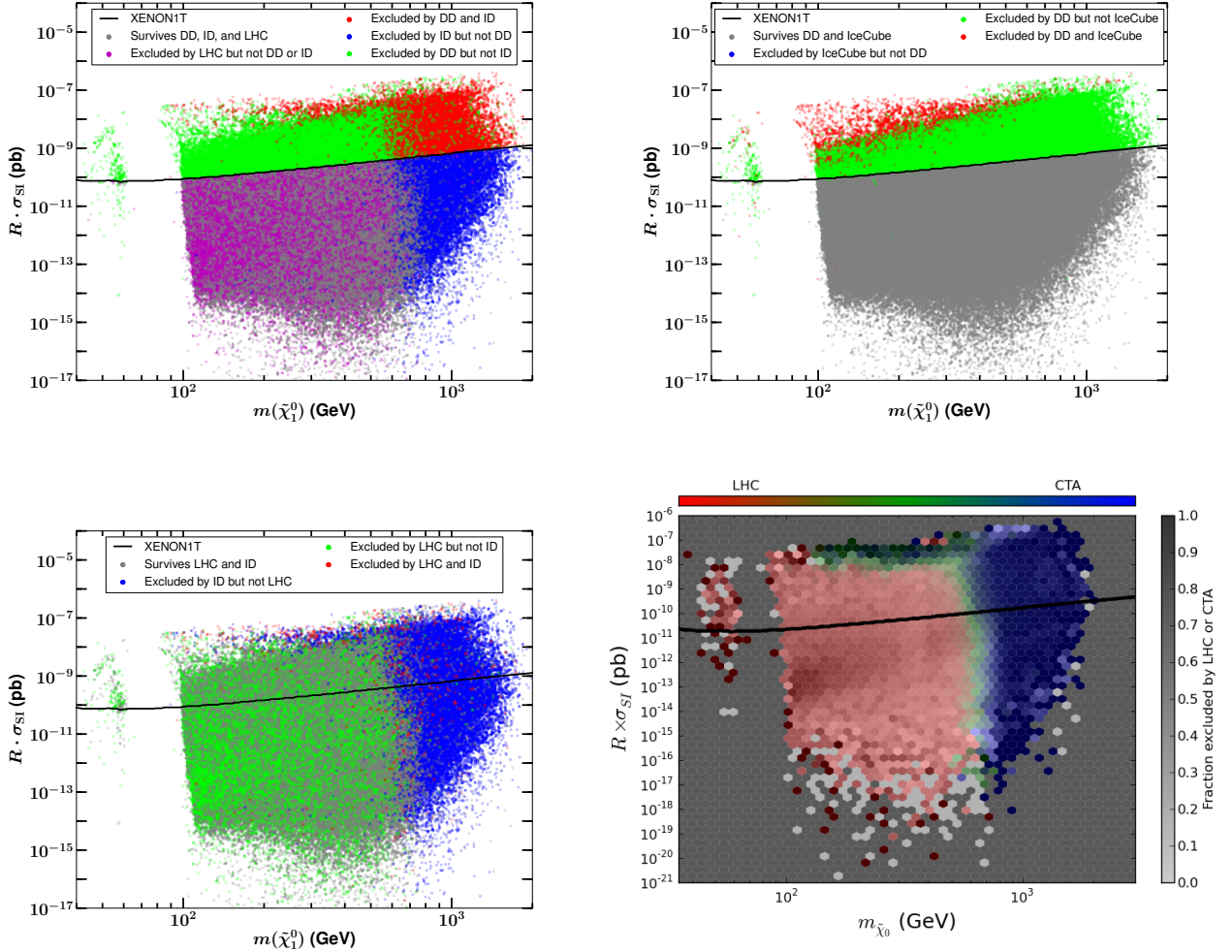


Figure 9: Comparisons of the models surviving or being excluded by the various searches in the LSP mass-scaled SI cross section plane as discussed in the text. The SI XENON1T line is shown as a guide to the eye.

Fig. 9 shows the survival and exclusion rates resulting from the various searches and their combinations in the LSP mass-scaled SI cross section plane. In the upper left panel we compare these for the combined direct detection (DD = XENON1T + COUPP500) and indirect detection (ID = *Fermi* + CTA) DM searches. Here we see that 11% (15%) of the models are excluded by ID but not DD (excluded by DD but not ID) while 8% are excluded

by both searches. On the other hand, we also see that 66% of the models survive both sets of DM searches; 41% of this subset of models, in turn, are excluded by the LHC searches. Note that the DD- and ID-excluded regions are all relatively well separated in terms of mass and cross section although there is some overlap between the sets of models excluded by the different experiments. In particular we see that the ID searches (here almost entirely CTA) are covering the heavy LSP region even in cases where the SI cross section is very low and likely beyond the reach of any potential DD experiment. Similarly, in the upper right panel we see that 22% (0%) of the models are excluded uniquely by DD (IceCube) only while 1% can be simultaneously excluded by both sets of searches and 77% would be missed by either search set. In the lower left panel, ID and LHC searches are compared and we see that 17% (32%) of the models would be excluded only by the ID (LHC) searches. However, 2% (50%) of the models are seen to be excluded by (or would survive) both searches. The strong complementarity between the LHC, CTA and XENON1T experiments is evident here as CTA probes the high LSP mass region very well where winos and Higgsinos dominate, XENON1T chops off the top of the distribution where the well-tempered neutralino LSP states dominate, while the LHC covers the relatively light LSP region (independent of LSP type) rather well. Of course the strength of the LHC coverage will significantly improve in the future as here only 7 and partial 8 TeV analysis results have been employed. In the lower right panel the relative contributions arising from the LHC and CTA searches to the model survival/exclusion are shown. Here the intensity of a given bin indicates the fraction of models excluded there by the combination of both CTA and the LHC while the color itself indicates the weighting from CTA (blue) and LHC (red). It is again quite clear that CTA completely dominates for large LSP masses (which correspond to the mostly wino and Higgsino LSPs) and also competes with the LHC throughout the band along the top of the distribution where the LSP thermal relic density is approximately saturating WMAP/Planck. The LHC covers the rest of the region but is not yet as effective as CTA (excluding a smaller fraction of models in the relevant bins). In the future we will examine how including the new 8 TeV LHC analyses and our extrapolation to 14 TeV will enhance the effectiveness of the LHC searches; we expect these improvements to be substantial.

In order to study the complementarity arising from the various searches in more detail, we found that it is very useful to project their individual capabilities onto several sets of fixed parameter planes that are directly related to one of the search categories. We begin these projections with Fig. 10 which compares the various search capabilities in the familiar LSP mass-scaled SI cross section plane, most relevant for the SI DD searches, as defined by the fraction of models excluded in the various bins. Here we see the contributions to the search coverage coming uniquely from CTA (top left), the LHC (top right), IceCube (bottom left) and *Fermi* (bottom right) and also show the limit that could arise from XENON1T for comparison. Here we see that both IceCube and *Fermi* probe models with low LSP masses and large SI cross sections, where the LSP tends to be a bino-Higgsino admixture; this region is also accessible to the DD experiments such as XENON1T. On the other hand, CTA has access to the heavy LSP region, where there are a large fraction of relatively pure winos and Higgsinos, while the LHC coverage is concentrated (for now) on the relatively low mass LSP

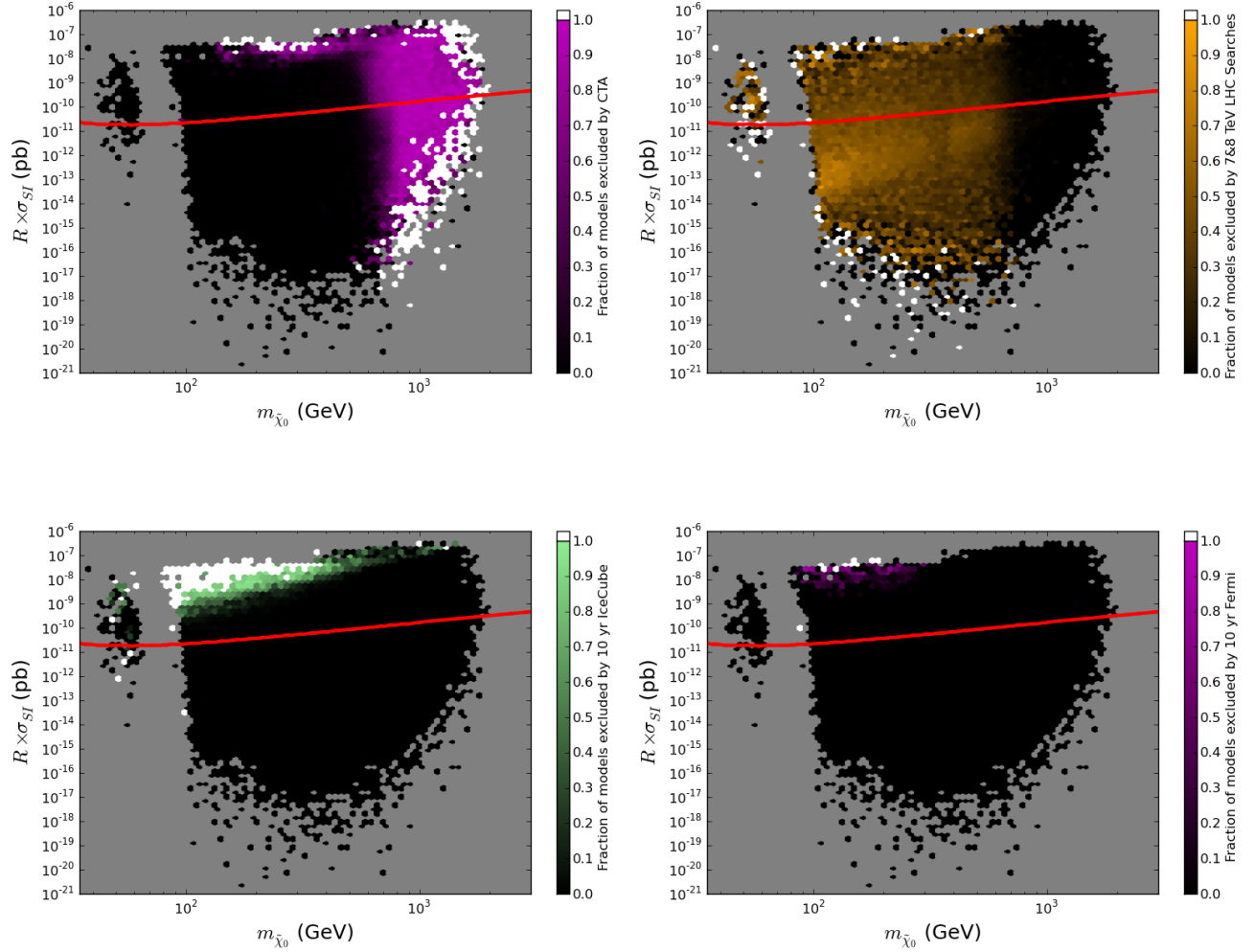


Figure 10: Comparisons of the various search capabilities in the LSP mass-scaled SI cross section plane as probed by the fraction of models excluded by CTA (top left), the LHC (top right), IceCube (bottom left) and *Fermi* (bottom right). The potential SI bound from XENON1T is also shown in all cases.

region. Interestingly, the LHC searches are not independent of the SI cross-section; a region of enhanced exclusion fraction is seen for SI cross sections near $\sim 10^{-13}$ pb. Models with cross sections in this region mostly have wino-like LSPs with light squarks; wino-like LSPs with heavier squarks have a lower SI cross-section, while Higgsinos and mixed states tend to have a higher SI cross-section whether or not light squarks are present.

Taking these same results we can now project them in to the plane which is most relevant for DM ID searches. Fig. 11 compares the various search capabilities as above but now in the LSP mass-LSP pair annihilation cross section plane, for which both *Fermi* and CTA are directly applicable. Here we see the fraction of models that would be excluded by searches at CTA (top left), the LHC (top right), IceCube (bottom left) and XENON1T (bottom right). The potential bounds from *Fermi* and CTA are also shown in all cases as curves penetrating through the upper left and right parts of the figures, respectively. Here the dashed (solid) curves correspond to the limit obtained when 100% of the LSP pair annihilates only into the $b\bar{b}$ (W^+W^-) final state; we emphasize that a generic LSP in a pMSSM model may annihilate to many different final states beyond these two cases. Note again that CTA is primarily sensitive to models with LSP mass above 100 GeV and with $\langle\sigma v\rangle$ relatively close to the thermal relic value. As in the previous figure, the LHC is mainly effective in the lower LSP mass region. In addition, the LHC searches are seen to be particularly efficient along a thin line starting at annihilation cross sections of $\sim 10^{-29} \frac{\text{cm}^3}{\text{s}}$ for 100 GeV LSPs and increasing with LSP mass; this line corresponds to models with a pure wino LSP. (Note that this line is below the projected reach of the ID searches for LSP masses accessible to the current LHC searches.) IceCube is seen to nicely complement CTA in the region with large cross sections but low LSP masses. The projection of the XENON1T coverage onto this plane is interesting with most of the exclusion appearing at lower LSP masses where the LSP pair annihilation cross section is also large. However, there is also an extended, somewhat diffuse region of substantial coverage throughout the entire right half of the parameter space as well as on the Z, h funnel ‘island’ at small LSP masses.

As a final example of these projections, we now consider the thermal relic density-LSP mass plane; the relic density plays a ‘linear’ role in the DD searches but a ‘quadratic’ one in the ID case (and none at all as far as the LHC is concerned). In Fig. 12 we display the coverage comparisons for the various experiments: CTA, LHC, IceCube, XENON1T and *Fermi*, respectively. Again we see that multiple searches can make overlapping contributions in various parts of this parameter space plane⁸. These plots show several things: First, CTA, as expected, does an excellent job at accessing most of the models with LSP masses above ~ 250 GeV that saturate the relic density. However, for larger masses the CTA coverage extends down to relic densities as much as a factor of ~ 10 or so below the WMAP/Planck value. *Fermi* is seen to cover only the low LSP mass region with relic densities not far from the thermal value while IceCube can go to much lower relic densities provided the LSP mass is below ~ 500 GeV or so. XENON1T displays coverage throughout this plane but does best for LSP masses below ~ 300 GeV, providing coverage even for models with very low relic

⁸In the case of an actual DM *discovery*, the existence of substantial overlapping regions in this parameter space will be of significant help in trying to determine the specific nature of the LSP.

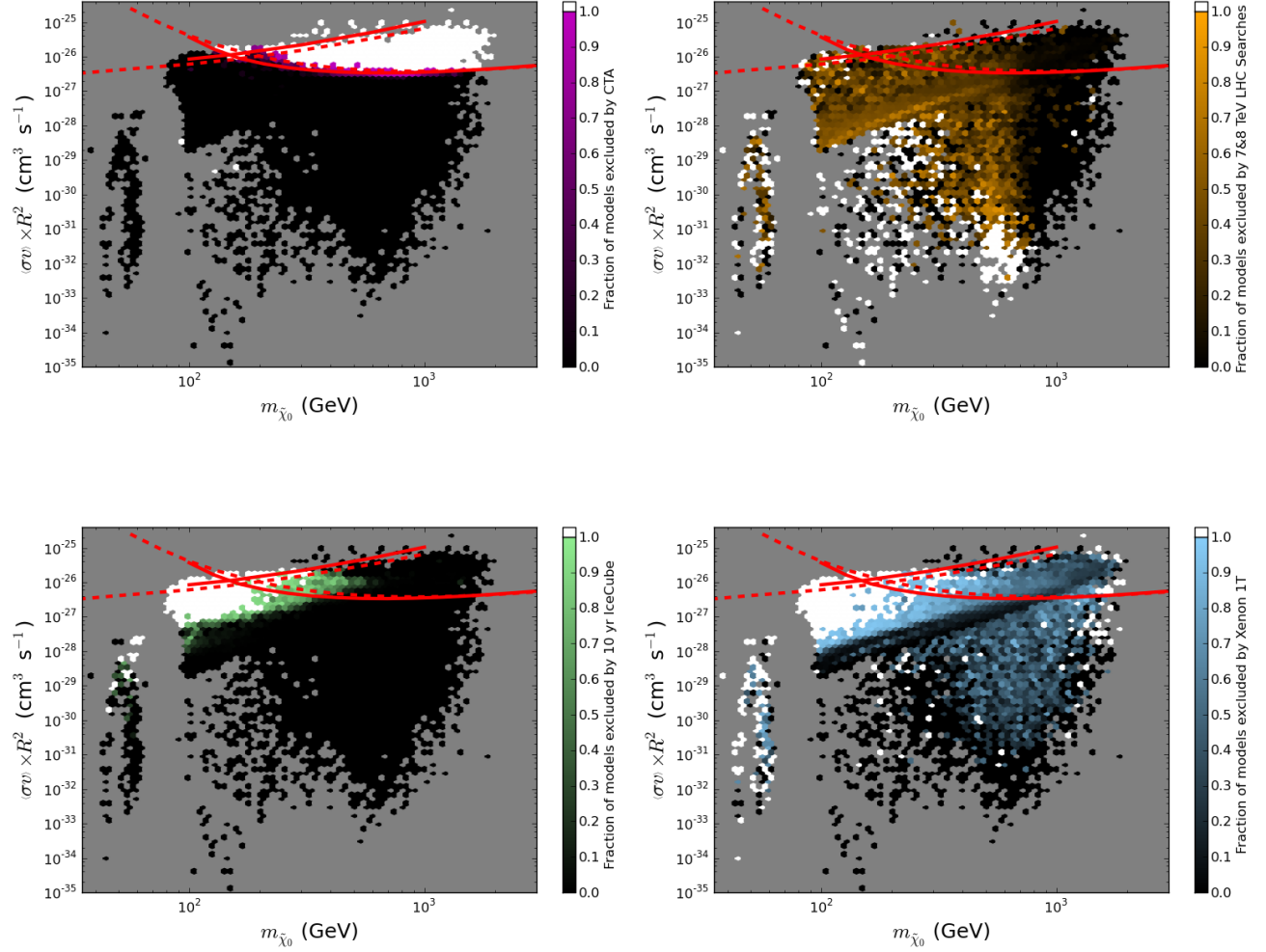


Figure 11: Comparisons of the various search capabilities in the LSP mass-LSP pair annihilation cross section plane, shown by the fraction of models in each bin which are excluded by CTA (top left), the LHC (top right), IceCube (bottom left) and XENON1T (bottom right). The potential bounds from *Fermi* and CTA are also shown in all cases as curves through the upper left and right parts of the figures, respectively. The dashed (solid) curves are for 100% annihilation into $b\bar{b}$ (W^+W^-) final states only.

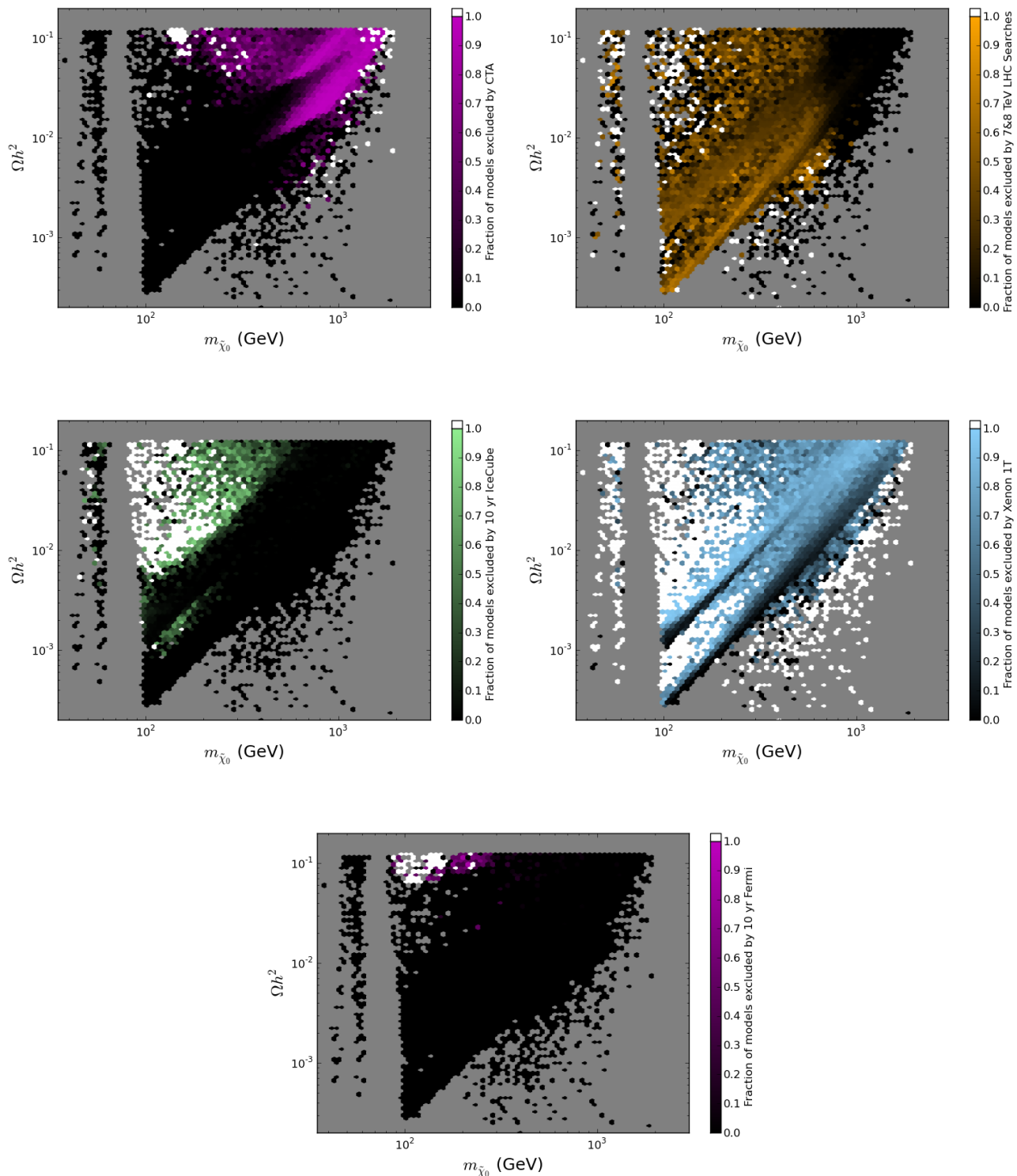


Figure 12: Comparisons of the various search capabilities in the LSP mass-relic density plane as probed by the fraction of models excluded by CTA (top left), the LHC (top right), IceCube (middle left), XENON1T (middle right) and *Fermi* (bottom).

densities. Of course, even for masses up to 1-2 TeV, XENON1T still provides quite decent model coverage in this parameter plane. As noted already, most of the impact of the LHC is at present seen to be at lower LSP masses below ~ 500 GeV. The LHC coverage is relatively uniform as far as the value of the relic density is concerned except in the case of very light LSPs where the coverage is very strong. Of course, we again remind the reader that we still need to add the additional information coming from the new 8 TeV LHC analyses not included here as well as the extrapolations to 14 TeV so that the coverage provided by the LHC should be expected to improve substantially.

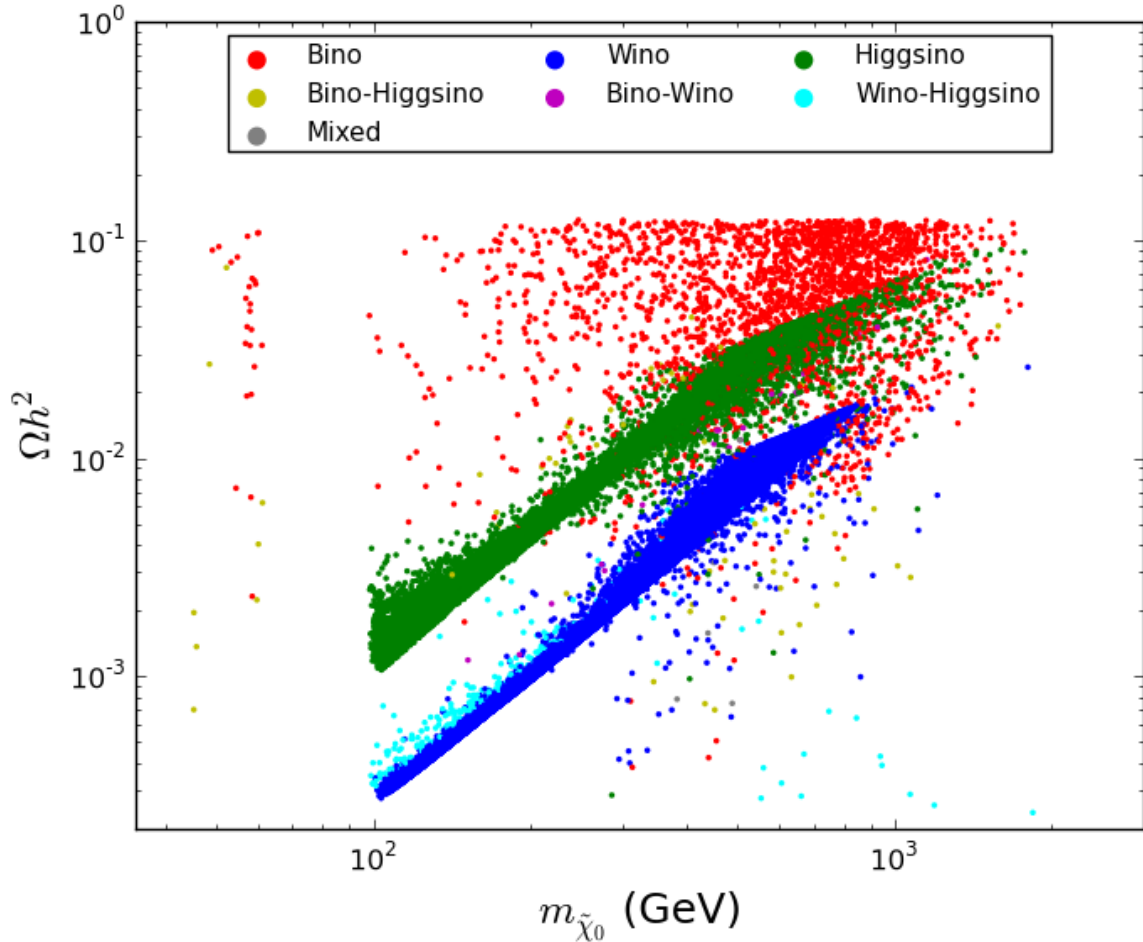


Figure 13: Thermal relic density as a function of the LSP mass for all pMSSM models, surviving after all searches, color-coded by the electroweak properties of the LSP. Compare with Fig. 2.

Finally, Fig. 13 shows the impact of combining all of the different searches in this same Ωh^2 -LSP mass plane which should be compared with that for the original model set as generated that is shown in Fig. 2. Here we see that (i) the models that were in the light h

and Z -funnel regions have essentially evaporated (and further measurements of the invisible width of Higgs as well as the standard SUSY searches will likely further restrict these), *(ii)* the well-tempered neutralinos are now seen to be completely gone, mostly due to XENON1T and IceCube; *(iii)* the possibility of almost pure Higgsino or wino LSPs even approximately saturating the relic density has vanished thanks to CTA and *(iv)* the mixed wino-Higgsino models, due to a combination of measurements, have also completely disappeared. *(v)* The only models remaining which *do* saturate the WMAP/Planck relic density are those with bino co-annihilation. *(vi)* We find that $\sim 61\%$ of all the pMSSM models have been excluded by at least one of the searches; this fraction should increase, possibly substantially, as more of the LHC searches are added to our existing analysis.

7 Conclusion

The results presented here are preliminary but directly lead us to a number of interesting conclusions which are already quite clear and which we expect to strengthen in the future:

- Even if the LSP *does not* make up all of the DM, it can still be observed in both direct and indirect detection experiments as well as in neutrino experiments such as IceCube. Of course, searches at the LHC are not influenced by the LSP relic density.
- The set of models remaining after all the searches are performed that saturate the thermal relic density consist only of those with (co)annihilating bino LSPs.
- SI direct detection, CTA, and the LHC do most of the heavy lifting in terms of complementary searches covering the pMSSM parameter space.
- Multiple/overlapping searches allow for extensive parameter space coverage which will be of particular importance if a DM signal is observed.
- Most of the experiments are seen to provide complementary probes of the pMSSM parameter space.
- The strength of the LHC component in these searches should increase significantly as further 8 TeV analyses as well as those from the 14 TeV LHC are included.

There is still much work remaining to be done, mostly related to the updating and extending of these pMSSM studies to include further LHC results, and this will be completed by the end of the Snowmass summer study.

In summary, the pMSSM provides an excellent tool for studying complementarity between different approaches to the search for dark matter.

Acknowledgments

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