

On the Origin of Light Dark Matter Species

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Abstract

TeV-mass dark matter charged under a new GeV-scale gauge force can explain electronic cosmic-ray anomalies. We propose that the CoGeNT and DAMA direct detection experiments are observing scattering of light stable states — “GeV-Matter” — that are charged under this force and constitute a small fraction of the dark matter halo. Dark higgsinos in a supersymmetric dark sector are natural candidates for GeV-Matter that scatter off protons with a universal cross-section of $5 \times 10^{-38} \text{ cm}^2$ and can naturally be split by 10–30 keV so that their dominant interaction with protons is down-scattering. As an example, down-scattering of an $\mathcal{O}(5)$ GeV dark higgsino can simultaneously explain the spectra observed by both CoGeNT and DAMA. The event rates in these experiments correspond to a GeV-Matter abundance of 0.2–1% of the halo mass density. This abundance can arise directly from thermal freeze-out at weak coupling, or from the late decay of an unstable TeV-scale WIMP. Our proposal can be tested by searches for exotics in the BaBar and Belle datasets.

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I. INTRODUCTION

For decades the WIMP — a particle with TeV-scale mass that freezes out through weak-strength couplings — has been the canonical dark matter candidate. The electronic cosmic ray excesses observed by PAMELA [1], Fermi [2], and HESS [3, 4], along with the discovery of the WMAP and Fermi “haze”[5–9], have both bolstered and complicated the WIMP picture. Together these experiments suggest that WIMP dark matter either annihilates or decays into lepton-rich final states (see e.g. [10]), with a rate 10-1000 times larger than expected from s -wave annihilation. Dark matter charged under a new GeV-scale $U(1)_D$ ‘dark force’ that kinetically mixes with hypercharge can explain these results through either Sommerfeld-enhanced annihilation [11–15] or decay [16] into light dark-sector states, which are in turn forced by kinematics to decay into leptons. Such a GeV-scale sector is naturally realized in models with supersymmetry-breaking near the weak scale, where D -term mixing dynamically generates a GeV-scale mass for the $U(1)_D$ gauge boson A' [17–21] (summarized in Section II).

The spontaneous breaking of this $U(1)_D$ gauge symmetry requires a higgs sector, and this sector can naturally contain stable GeV-scale states such as dark higgsinos. These and other stable “GeV-Matter” particles charged under $U(1)_D$ interact with protons with a cross-section

$$\sigma_{\chi,p} \approx 5 \times 10^{-38} \text{ cm}^2 \tag{1}$$

that is **completely fixed by Standard Model couplings** [18]. Because this cross section is so large, even a small abundance of GeV-Matter can be observed at direct detection experiments. In fact, we should expect GeV-Matter states to comprise only a small fraction of the dark matter halo because their annihilation cross-section is large ($\sim \alpha^2/m_{\text{GeV}}^2$).

Two experiments have reported anomalous events that may arise from scattering of a light dark matter species. DAMA [22–25] has reported a 9σ annual modulation signal at 2–5 keVee in their NaI crystal detector. Early this year, the CoGeNT collaboration reported about 100 events from an unknown source in a low-threshold, high-resolution Ge detector [26]; these events are consistent with light dark matter scattering with a mass and cross-section similar to those needed to explain the DAMA modulation. As shown in Figure 1 (see Section V for details), both signals can be explained by 3–12 GeV dark matter scattering with a cross-section $\sigma_{\text{eff}} \sim (1 - 5) \times 10^{-40} \text{ cm}^2$, or by a fraction of the halo scattering with

a larger cross-section (see also [26–28]). If the events reported by these experiments arise from light dark matter scattering, they provide further evidence for a new sector at the GeV scale.

In light of the cross section (1), dark sector higgsinos at 3-12 GeV naturally explain the CoGeNT and DAMA excesses if they comprise a fraction

$$f_{\text{light}} \equiv \frac{\rho_{\text{light}}}{\rho_{\text{DM}}} \approx (2-10) \times 10^{-3} \quad (2)$$

of the halo mass density. We propose two natural origins for an abundance of this size:

Thermal freeze-out generates an abundance (assuming s -wave annihilation)

$$f_{\text{light}} \sim 3 \times 10^{-3} \left(\frac{m_{\text{light}}}{6 \text{ GeV}} \right)^2 \left(\frac{1/150}{\alpha} \right)^2, \quad (3)$$

which yields the desired GeV-Matter relic density for small couplings, as we discuss in Section III.

Late Decays of TeV-mass WIMP-sector states generate a density

$$\frac{n_{\text{light}}}{n_{\text{heavy}}} \sim 1, \quad f_{\text{light}} \sim \frac{m_{\text{light}}}{m_{\text{heavy}}} \frac{\rho_{\text{heavy}}}{\rho_{\text{DM}}} \sim \text{few} \times 10^{-3} \quad (4)$$

of GeV-Matter that can easily dominate over its thermal abundance. This sharp prediction is insensitive to the details of the decay process, and agrees with the halo fraction required to explain the CoGeNT and DAMA signals! The decaying TeV-mass particle can be the scalar superpartner of a stable fermionic WIMP, as discussed in Section IV.

We therefore propose a new unified framework for interpreting both the direct detection and cosmic ray anomalies from the same dynamics.

$U(1)_D$ -breaking mass splittings of $\mathcal{O}(\text{few keV} - \text{few GeV})$ for both WIMPs and GeV-Matter arise from generic interactions with TeV-scale states. In the presence of such mass splittings, dark matter scatters through inelastic up- and down-scattering, with elastic processes highly suppressed. These splittings have two important effects: when the GeV-Matter states are split by $\mathcal{O}(10 \text{ keV})$, direct detection is dominated by down-scattering, which improves the agreement between DAMA and CoGeNT and permits lower masses of the light dark matter than in the elastic case, as discussed in Section V and more generally in [29].

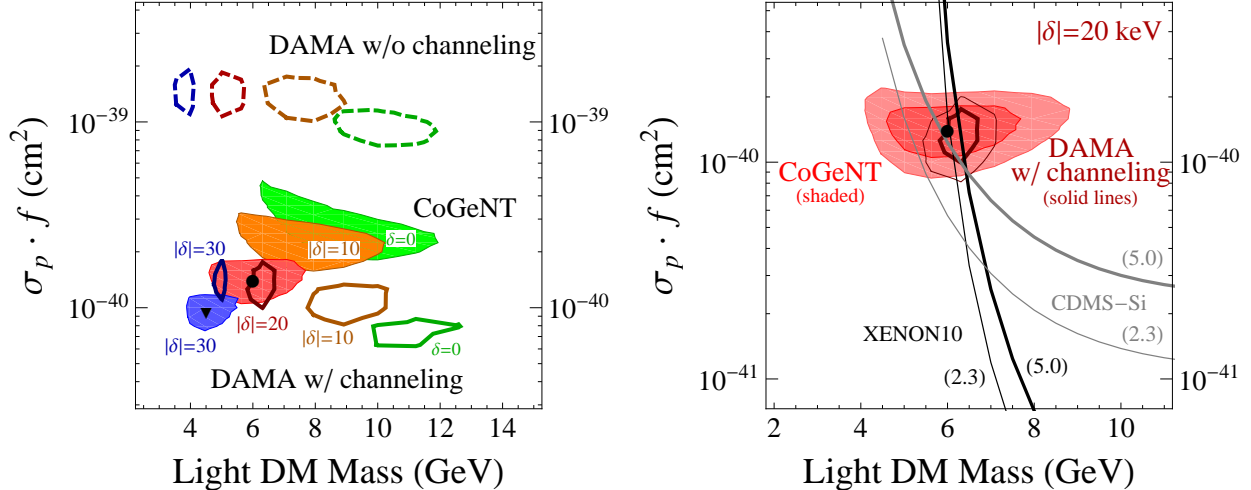


FIG. 1: **Left:** Regions of scattering cross-section per nucleon times GeV-Matter mass fraction in the halo versus mass favored by CoGeNT (filled regions), DAMA with channeling (solid-line contours), and DAMA without channeling (dashed contours) for light dark matter scattering coherently off protons. We show favored regions (from right to left) for elastic scattering (green) and inelastic splittings $|\delta| = 10$ keV (orange), 20 keV (red), and 30 keV (blue), with equal fractions of up- and down-scattering. The contours correspond to “ 1.6σ ,” neglecting systematic uncertainties. The black dot and upside-down triangle correspond to benchmark points we consider in Figure 2. **Right:** A close-up of the $|\delta| = 20$ keV allowed region with “ 3σ ” contours added and constraints from XENON10 (black lines) and CDMS-Si (gray lines) computed for $|\delta| = 20$ keV. Thick (thin) lines denote 5.0 (2.3) events expected, again neglecting systematics. More details are given in §V.

A larger \gtrsim MeV splitting of the TeV-scale WIMP allows it to explain the INTEGRAL 511 keV excess and renders it practically invisible to direct detection experiments [11, 30].

The success of these models is particularly striking in light of the difficulty of generating such high direct detection rates through Standard Model processes, even when a large abundance of light dark matter is generated through an asymmetry [31]. Dark matter coupled to the Z boson can explain the reported rate, but is quite constrained by LEP measurements of the Z invisible width [32]. Standard Model Higgs exchange, in contrast, leads to typical cross-sections $\lesssim 10^{-43}$ cm² (see e.g. [33]) because the Higgs couples only weakly to nuclei. Explanations of the CoGeNT signal through scalar interactions posit a new light scalar that mixes significantly with the Higgs, but remains relatively light through a percent-level fine-

tuning [27, 28]. In contrast, the low mass of our dark matter and mediator arise from simple dynamics.

Implications for B -factories

Analysis of existing low energy collider data is crucially important for testing the GeV-Matter scenarios in this paper (see e.g. [34–39]). The strongest constraint on $U(1)_D$ gauge boson decays to Standard Model matter is from a BaBar search for a resonance in $\gamma\mu^+\mu^-$ in $\Upsilon(3s)$ data [40], which is sensitive to the “radiative return” process $e^+e^- \rightarrow \gamma A'$. The parameter space consistent with the GeV-Matter explanation of the CoGeNT and DAMA excesses is just below the limit from this search, but **can be tested by a search in the full Belle and BaBar datasets**, which together comprise about a factor of 30 higher statistics than the $\Upsilon(3s)$ data. The “higgs'-strahlung” process [35] may also be accessible at B -factories, leading to a striking six-lepton final state. Searches at J/ψ - and ϕ -factories can also explore the low-mass parameter space [36, 37, 41, 42].

Down-Scattering and Lighter Dark Matter

Three flaws have been noted with elastically scattering light dark matter as an explanation of the DAMA annual modulation signal: the predicted modulation spectrum grows exponentially at low energies while the DAMA spectrum appears to fall near threshold [43]. Moreover, the mass range favored by both the DAMA and CoGeNT spectra, ~ 10 GeV, is disfavored by XENON10 [44–46] and especially CDMS Silicon [47] data (see [27]). Finally, the scattering rates required to explain the CoGeNT and DAMA signals differ by about a factor of 5 (though this final discrepancy depends on the precise channeling properties of Iodine [27, 48–50]).

However, there is no reason to expect the light dark matter scattering to be purely elastic. Mass splittings of order 10-30 keV for GeV-Matter are readily generated by TeV-suppressed higher-dimension operators.

Because the dark sector kinetically decouples from the Standard Model at temperatures high compared to this splitting, both states will be equally populated and the heavy state will be cosmologically long-lived [51–53]. The excited states interact with nuclei through

down-scattering, which pushes the nuclear recoil spectrum to higher energies for a given dark matter mass. As we discuss in §V, this effect improves the agreement of light dark matter with DAMA modulation data (relative to the elastic case), brings the CoGeNT and DAMA expected cross-sections together, and is consistent with the XENON10 null search. The residual tension of this scenario with the 5-tower CDMS Silicon analysis [47] is extremely sensitive to the experimental selection efficiency near threshold and small energy-scale systematics. There is also tension with the DAMA unmodulated spectrum in [24]. A general analysis of down-scattering light dark matter [29] also finds a lower-mass region consistent with DAMA and CDMS Silicon results.

II. AN APPROXIMATELY SUPERSYMMETRIC DARK SECTOR

Our starting point is a simple extension ([18–21]) of the supersymmetric $SU(3) \times SU(2)_L \times U(1)_Y$ Standard Model that dynamically generates a mass scale of $\mathcal{O}(\text{GeV})$ for the “dark force.” The dark sector consists of two chiral superfields h_u and h_d that are oppositely charged (± 1) under a new $U(1)_D$ gauge group, but neutral under the Standard Model gauge groups. No Standard Model fields are directly charged under $U(1)_D$. We also add a gauge singlet S , and consider the following superpotential for the GeV-scale sector:

$$\mathcal{W} = \lambda S h_u h_d - \frac{1}{4} W_D^2 - \frac{\epsilon}{2} W_Y^\alpha W_{\alpha D}. \quad (5)$$

Here, W_D and W_Y are the $U(1)_D$ and hypercharge vector superfields, and the last term is the supersymmetric version of the usual gauge kinetic mixing,

$$\mathcal{L} = \frac{\epsilon}{2} F_{Y,\mu\nu} F_D^{\mu\nu}. \quad (6)$$

Besides inducing the gauge kinetic mixing (6) and gaugino kinetic mixing, the last term also induces a fixed ϵ -suppressed potential coupling of Standard Model and dark-sector scalar fields through the mixing of Standard Model and dark-sector D -terms. After electroweak symmetry breaking, this mixing generates a dark-sector scalar potential

$$V = \lambda^2 (|S h_u|^2 + |S h_d|^2 + |h_u h_d|^2) + \frac{g_D^2}{2} (|h_u|^2 - |h_d|^2 + \zeta^2)^2, \quad (7)$$

that includes a negative mass-squared for the scalar h_d (or h_u if $\zeta^2 < 0$), where

$$\zeta^2 = \frac{\epsilon g_Y |\cos 2\beta|}{2g_D g_2^2} m_W^2 \quad (8)$$

is the hypercharge D -term vev.

The stable vacuum is one where $U(1)_D$ is broken but supersymmetry is approximately preserved (supersymmetry-breaking corrections, discussed below, are higher-order in ϵ). The light fields are grouped into massive gauge and chiral supermultiplets, with masses

$$m_{A'} = m_{\tilde{\lambda}} = \sqrt{2}g_D\zeta \quad \text{and} \quad m_h = m_{\tilde{h}} = \lambda\zeta. \quad (9)$$

The uneaten real part of h_d and \tilde{h}_d (which gives a Dirac mass to the gaugino $\tilde{\lambda}$) are in the massive gauge supermultiplet, while the h_u and S scalars and their superpartners \tilde{h}_u and \tilde{S} marry into the massive chiral supermultiplet. The A' mediates scattering of particles with $U(1)_D$ -charge q_d off Standard Model matter **with a non-relativistic scattering cross-section almost completely determined by Standard Model parameters**,

$$\sigma_{\chi,p} = 16\pi q_d^2 \frac{\mu_{\chi,p}^2 \alpha_D \alpha c_W^2 \epsilon^2}{m_{A'}^4} = \frac{(16\pi q_d^2 \alpha_2^2 c_W^4) \mu_{\chi,p}^2}{(\cos 2\beta)^2 m_W^4} \approx 5 \times 10^{-38} \left(\frac{2q_d}{\cos 2\beta} \right)^2 \text{ cm}^2, \quad (10)$$

with all dependence on ϵ and the dark gauge coupling g_D canceling and only mild dependence on the ratio $\tan \beta$ of Standard Model Higgs vevs and on the dark matter mass (we chose $m_\chi = 5$ GeV in the numerical estimate above).

In addition to these light fields, we assume throughout that the dominant WIMP component of dark matter arises from a vector-like pair of chiral superfields $\Phi_{u,d}$ with charge $\pm q$, so that the full renormalizable superpotential is

$$\mathcal{W}_{\text{tot}} = \lambda S h_u h_d - \frac{1}{4} W_D^2 - \frac{\epsilon}{2} W_Y^\alpha W_{\alpha D} + M \Phi_u \Phi_d, \quad (11)$$

with M at the TeV-scale (for example, M and the μ -term of the MSSM Higgs fields may have the same origin). The stability of Φ and absence of other couplings to the light fields can easily be guaranteed by global symmetries, in particular including a \mathbb{Z}_2 under which the $\Phi_{u,d}$ fields are odd. Alternately, the stability of Φ is automatic if its charge q is non-integer.

Supersymmetry-Breaking Effects

Supersymmetry-breaking effects are crucial for dark sector physics; we will always take fermionic states as our dark matter candidates, so we will insist that they are stable. In the light GeV-Matter sector, some scalar density is innocuous but it is important that the Φ scalars decay.

We begin by discussing the supersymmetry-breaking effects that arise inevitably from gauge mediation through the kinetic mixing and from anomaly mediation, and then we will consider further effects that can arise through gravity or direct mediation to the dark sector.

In models of gauge mediation, all scalars charged under $U(1)_D$ receive soft masses mediated from the messengers through kinetic mixing of order [20, 54]

$$\tilde{m}_{GMSB}^2 \sim \epsilon^2 \frac{\alpha_D}{\alpha_Y} m_{E^c}^2, \quad (12)$$

where m_{E^c} is the mass of the right-handed selectron in the Minimal Supersymmetric Standard Model (MSSM). A *negative* supersymmetry breaking mass-squared for S is generated by RG evolution [20], although it is suppressed by an additional loop factor times a potentially large logarithm. The R -breaking $Ah_u h_d S$ and $B\Phi_u \Phi_d$ terms are suppressed by $\mathcal{O}(\alpha_D \epsilon^2)$ relative to the MSSM bino mass, $m_{\tilde{B}}$. A Majorana soft mass for the dark gaugino is given by $\epsilon^2 m_{\tilde{B}}$, while the MSSM bino and dark gaugino have a mixing mass of $\epsilon m_{\tilde{B}}$. In the absence of $U(1)_D$ symmetry-breaking, these would preserve a zero mass eigenvalue. When $U(1)_D$ is broken they only induce a tiny splitting $\epsilon^2 m_{A'}/m_{\tilde{B}}$ of the Dirac gauginos into two Majorana states (approximately $(\tilde{\lambda} \pm \tilde{h}_d)/\sqrt{2}$), one heavier and the other lighter than the A' . Both the lighter and heavier of these states can decay through kinetic mixing with the MSSM bino to $\gamma \tilde{G}$, with a lifetime [20]

$$\tau_{\tilde{\lambda}} \sim 10^3 \text{ s} \left(\frac{\sqrt{F}}{10^5 \text{ GeV}} \right)^4 \left(\frac{1 \text{ GeV}}{m_{\tilde{\lambda}}} \right)^5 \left(\frac{10^{-3}}{\epsilon} \right)^2 \quad (13)$$

as long as $m_{\tilde{\lambda}} > m_{3/2}$. Constraints from BBN [55] and from EGRET searches for γ -ray lines from the galactic center require [56] either very low $\sqrt{F} \lesssim 10^5$ GeV for prompt gaugino decays or rather high $\sqrt{F} \sim 10^{10}$ GeV, so that the gravitino is heavier than the gaugino and the decay is forbidden. In the latter case, the dark sector soft masses inevitably receive anomaly-mediated contributions [57, 58] of order

$$\tilde{m}_{AMSB} \sim \frac{\alpha_D}{4\pi} m_{3/2} \sim 1 - 100 \text{ MeV}, \quad (14)$$

which are comparable to the gauge-mediated contributions to the scalar soft masses and dominate the gaugino soft mass.

If F is low, dark matter annihilation or decay to A' 's at a rate compatible with the FERMI e^+e^- excess could potentially generate a gamma-ray signal visible with existing

or upcoming satellite data. This is because supersymmetric couplings will also allow dark matter annihilation/decay to $\tilde{\lambda}$, whose subsequent decay produces a photon. This signal is fairly model-dependent — it is reduced if dark matter annihilates or decays through supersymmetry breaking couplings, or if the gaugino decays primarily through R -parity-violating interactions. Nonetheless, it would be interesting to investigate these signals in future work.

Although the effects we considered above are irreducible, it is plausible to include additional supersymmetry breaking operators such as

$$\int d^4\theta \frac{X^\dagger X}{M_*^2} S^\dagger S \quad \text{and} \quad \int d^2\theta \frac{X}{M_*} W_D^2, \quad (15)$$

where X is a supersymmetry-breaking-spurion with an F_X component vev. We can think of these terms as originating from ‘gravity mediation’ with $M_* \sim M_{pl}$, but it is equally natural for M_* to be much smaller. These operators can contribute to the Majorana gaugino mass and also raise the scalar higgs sector (S and h_u) soft masses. In particular, the scalar S can obtain a net positive supersymmetry-breaking-mass, forbidding the decay of the lightest higgsino to S and a gravitino.

Stable States and Their Properties

We will always take our GeV-Matter (the light stable states) to be fermionic, so in this section we will be interested in the properties of these fermions and in the mechanisms through which we avoid stable scalars.

The Dirac gaugino *can* be stabilized on cosmological timescales if the supersymmetry-breaking F is large, so that either the lifetime ($\propto F^2$) of the gaugino is very large or $m_{3/2} > m_{\tilde{\lambda}}$ (which requires $\sqrt{F} \gtrsim 0.5 \times 10^{10}$ GeV) so that it is kinematically forbidden to decay into the gravitino. Nonetheless, it is not a good candidate to explain the CoGeNT and DAMA signals in our scenario. For such high F , the gaugino anomaly-mediated soft mass $\sim 10 - 100$ MeV is sufficiently large that the heavier gaugino decays rapidly to e^+e^- and the lighter gaugino. Though the lighter component can comprise a non-negligible fraction of the dark matter, it can only scatter elastically (because the excited state is depopulated and up-scattering is kinematically forbidden) with a cross-section that is velocity-suppressed and thus too small to explain the direct detection signals.

The higgs supermultiplet thus contains the only viable GeV-Matter candidates within the model (5) that can explain CoGeNT and DAMA. We assume that the higgs supermultiplet is stabilized on cosmological timescales by a discrete or continuous global symmetry. We are primarily interested in higgsino GeV-Matter $((\tilde{h}_u \pm \tilde{S})/\sqrt{2})$, which readily acquires a small inelastic splitting as we discuss below. The higgsinos are stable under either of two conditions: either S must receive a positive soft-mass contribution that makes it heavier than the higgsino, or F must be sufficiently large that decays $\tilde{h} \rightarrow S + \tilde{G}$ are kinematically forbidden.

If F is low and the supersymmetry breaking contribution to the S scalar mass is positive (which requires some contribution besides gauge mediation), all scalars can decay to their fermionic partners and a gravitino with a lifetime given parametrically by $\tau \sim 16\pi F^2/(m_h^4 \Delta m)$, and likewise gauginos can decay to $\gamma \tilde{G}$ according to (13). The higgsino is then the *only* stable relic in the light dark sector. In this case, the scalar $\tilde{\Phi}$ also typically decays to $\tilde{\Phi} + \tilde{G}$.

Large F also leads to a stable higgsino if the gravitino mass is large enough to kinematically forbid higgsino decays to $S + \tilde{G}$ or $S + \tilde{h}_d$. In this scenario, one or both scalars in the higgs multiplet may be stable, but they are less phenomenologically important than the higgsinos. In particular, S is neutral under $U(1)_D$ and therefore only has a very small direct detection cross-section induced by its mass mixing with h_u . Even at high F , a heavier h_u scalar can typically decay through the same mass mixing to S and an off-shell A' . Even if the h_u is cosmologically stable, it typically has no inelastic splitting, making it less observable in direct detection than the higgsinos which, as we will discuss below, typically down-scatter.

A final possibility that we will not explore further in this work is that the higgsinos are actually lighter than the gaugino and gravitino states, so that the higgsinos are stabilized by R-parity.

Although supersymmetry-breaking effects leave the higgsino and $\tilde{\Phi}$ as exactly degenerate Dirac fermions (in contrast with the gaugino), interactions with 1-10 TeV states can lead to splittings by $\mathcal{O}(10 \text{ keV} - 1 \text{ MeV})$ when these heavy states are integrated out, for example, through the higher-dimension superpotential operator

$$\frac{1}{M_*} h_d^2 h_u^2. \quad (16)$$

Elastic scattering processes will be velocity suppressed and therefore negligible in all cases,

so the higgsinos will only be visible to direct detection experiments through inelastic down-scattering.

In many models (see Appendix B), the same interactions generate an $\mathcal{O}(\text{MeV})$ or larger splitting for the heavy dark matter fermions $\tilde{\Phi}_{u,d}$ and a smaller $\mathcal{O}(10 \text{ keV})$ splitting for the higgsinos; the simplest example is to take $q = 1/2$ and include the marginal superpotential operator $h_u \Phi_d^2$. For splittings $\delta M \gtrsim 1 \text{ MeV}$, the excited state of the TeV-scale dark matter decays to the lower-mass state and an e^+e^- pair, depopulating the initial state of the down-scattering process, and up-scattering off nuclei is kinematically forbidden. In this case, the leading direct detection cross-section would be elastic, but it is suppressed by both $(\frac{\delta M}{M})^2 \sim 10^{-12} - 10^{-6}$ and by a factor of the non-relativistic velocity squared $\sim 10^{-6}$, making it completely negligible. For splittings $\sim 1 \text{ MeV}$ the $\tilde{\Phi}_{u,d}$ self-interactions may explain the INTEGRAL 511 keV excess[11].

We note that it is less generic in these sectors to generate inelastic splittings for the scalar Φ than for the $\tilde{\Phi}$ fermion, so we insist that the scalar Φ decay. At low F the gauge-mediated splittings are sufficient to allow decays $\Phi \rightarrow \tilde{\Phi} + \tilde{G}$. For the higher- F scenario and heavier gravitinos, it is plausible that the Φ couples more directly to SUSY-breaking than do the fields in the low-mass dark sector (for example, if its μ -term is related to SUSY-breaking), so that decays to the gravitino are still kinematically allowed. Alternately, in Section IV we consider interactions with an intermediate-scale supersymmetry-breaking sector that lead to Φ scalar decays but keep the fermion $\tilde{\Phi}$ cosmologically stable.

III. GEV-MATTER FROM THERMAL FREEZE-OUT

Given the proximity of the CoGeNT and DAMA mass scales to the GeV-scale, a natural possibility is that the higgsinos of the dark sector are themselves stable and give rise to observable direct detection signals. We focus on $\lambda \gtrsim \sqrt{2}g_D$, so that the higgs supermultiplet is heavier than the gauge supermultiplet, and we will assume that global symmetries guarantee that this multiplet is stable on time-scales of order the age of the universe. Note that the gaugino states could be cosmologically stable, but they are not a viable candidate to explain the CoGeNT and DAMA signals as we discussed in Section II. The higgsino relic abundance agrees with (2) in a small-coupling limit. For larger couplings the thermal higgsino abundance is small but, as we will see in the next section, the late decay of a TeV-mass

WIMP gives rise to a relic density of higgsinos of the appropriate size.

We will also discuss how an additional light vector-like species (not part of the dark gauge or higgs sector) is a viable GeV-matter candidate if it is charged under $U(1)_D$.

Dark Higgsino Freeze-Out

Higgsino annihilation during freeze-out is dominated by annihilation to the real scalar h_d and a longitudinal A' , with cross-section

$$\sigma_{\tilde{h}} = \frac{5\pi\alpha_\lambda^2}{8m_{\tilde{h}}^2 v} \quad (17)$$

where $\alpha_\lambda = \lambda^2/(4\pi)$. This dominates over all other annihilation modes, which are either p -wave suppressed or suppressed by g^2/λ^2 , leading to an abundance

$$f_{\tilde{h}} \equiv \frac{\Omega_{\tilde{h}}}{\Omega_{DM}} \simeq 3.0 \times 10^{-3} \left(\frac{1/150}{\alpha_\lambda} \right)^2 \left(\frac{m_{\tilde{h}}}{6 \text{ GeV}} \right)^2, \quad (18)$$

where $\alpha_\lambda = \lambda^2/(4\pi)$, and we have chosen the baseline $m_{\tilde{h}} = 6 \text{ GeV}$ at the intersection of the CoGeNT and DAMA best-fit regions for 20 keV inelastic splitting (see Section V). Using (10) with $q_d = 1/2$, this gives rise to an effective direct-detection cross-section

$$\begin{aligned} \sigma_{\text{eff}} \equiv f_{\tilde{h}} \sigma_{\tilde{h},p} &\simeq 1.5 \times 10^{-40} \text{ cm}^2 \\ &\times \left(\frac{1}{\cos 2\beta} \right)^2 \left(\frac{1/150}{\alpha_\lambda} \right)^2 \left(\frac{m_{\tilde{h}}}{6 \text{ GeV}} \right)^2, \end{aligned} \quad (19)$$

Normalizing this to the σ_{eff} derived from DAMA and CoGeNT data determines the dark-sector higgs vev within this setup (at around 20 GeV), and the gauge coupling as a function of the A' mass:

$$\alpha_D = \frac{1}{2} \alpha_\lambda \left(\frac{m_{A'}}{m_{\tilde{h}}} \right)^2 \simeq 10^{-4} \left(\frac{m_{A'}}{1 \text{ GeV}} \right)^2 \left(\frac{m_{\tilde{h}}}{6 \text{ GeV}} \right)^{-1} \left(\frac{\sigma_{\text{eff}}}{1.5 \times 10^{-40} \text{ cm}^2} \right)^{-1/4} \quad (20)$$

for the CoGeNT/DAMA best fit. A separation of scales between the A' and higgsino masses, as suggested by the lepton-rich cosmic ray signals, implies a rather small dark gauge coupling in this scenario. This in turn can be translated into a prediction for ϵ using (8) and (9),

$$(\epsilon c_W)^2 \simeq 1.9 \times 10^{-5} \times \left(\frac{m_{A'}}{1 \text{ GeV}} \right)^2 \left(\frac{m_{\tilde{h}}}{6 \text{ GeV}} \right)^{1/2} \left(\frac{\sigma_{\text{eff}}}{1.5 \times 10^{-40} \text{ cm}^2} \right)^{1/4}. \quad (21)$$

If the A' decays directly to e^+e^- , models that produce σ_{eff} as a thermal relic are very near the sensitivity limit of existing B -factory searches discussed in Section VI.

For larger gauge couplings, the relic abundance of the higgsinos from thermal freeze-out is too small to explain the CoGeNT and DAMA signals. However, as we will discuss in Section IV, a relic abundance of the desired size can easily be generated by other means.

Vector-Like GeV-Matter

Now let us consider a slightly less minimal model that exhibits the parametrics of equation (3) without a direct link between the GeV-Matter mass and the A' mass. Consider a light vector-like species χ_{\pm} with superpotential $W \supset \mu_{\chi}\chi_{+}\chi_{-}$ and $\mu_{\chi} \sim 5\text{-}10$ GeV. Like $\Phi_{u,d}$, the χ_{\pm} states may be stabilized by a discrete symmetry such as a Z_2 where $\chi_{\pm} \rightarrow -\chi_{\pm}$. A higher-dimension coupling of χ to the higgs-sector field h_d such as

$$\frac{1}{\Lambda}\chi_{+}^2 h_d^2 \quad (22)$$

will lead to a ~ 10 keV splitting for the fermionic light dark matter states $\tilde{\chi}_{\pm}$ when $\langle h_d \rangle \sim 0.1\text{--}1$ GeV and $\Lambda \gtrsim 1\text{--}10$ TeV. Note that the scalar states in the χ multiplets can decay to $\tilde{\chi} + \tilde{G}$ or $\tilde{\chi} + \tilde{\lambda}$ pairs depending on the gravitino and gauge supermultiplet masses.

In a minimal scenario, the χ relic abundance is set by the annihilation cross-section to lighter higgs and gauge-multiplet states, which we parametrize as

$$\sigma_{\tilde{\chi}} \simeq \frac{3\pi\alpha_D^2}{2m_{\chi}^2 v} \quad (23)$$

This gives rise to a relic density

$$f_{\tilde{\chi}} \equiv \frac{\Omega_{\tilde{\chi}}}{\Omega_{DM}} \simeq 6 \times 10^{-4} \left(\frac{1/100}{\alpha_D} \right)^2 \left(\frac{m_{\chi}}{6 \text{ GeV}} \right)^2. \quad (24)$$

The χ scatter in direct detection experiments with total cross-section (10) with $q_d = 1$, giving rise to an effective cross-section

$$\begin{aligned} \sigma_{\text{eff}} \equiv f_{\chi} \frac{\Omega_{\tilde{\chi}}}{\Omega_{DM}} \sigma_{\tilde{h},p} &\simeq 1.2 \times 10^{-40} \text{ cm}^2 \\ &\times \left(\frac{1}{\cos 2\beta} \right)^2 \left(\frac{1/100}{\alpha_D} \right)^2 \left(\frac{m_{\chi}}{6 \text{ GeV}} \right)^2. \end{aligned} \quad (25)$$

The same parametrics applies to the heavy $\Phi_{u,d}$ (again, assuming that annihilation through gauge interactions dominates), albeit with a higher g_* in the freeze-out calculation. For example, $\alpha_D = 1/100$ used above would lead to $\Omega_{\Phi} = \Omega_{DM}$ for $\mu_{\Phi} = 550$ GeV. The above model is similar to that discussed in [59], but we consider a different hierarchy of masses and couplings.

IV. GEV-MATTER FROM THE DECAY OF HEAVY WIMPS

In this section, we will present a simple mechanism that can simultaneously explain the cosmic ray data and the direct detection signals. Our mechanism populates the GeV-Matter states via the late decay of a TeV-scale WIMP-sector particle, resulting in a relic abundance that precisely fits the CoGeNT and DAMA direct detection rates (with the scattering cross-section in Eq. (10)); it explains the cosmic-ray data through the decay of a TeV-scale WIMP with a lifetime on the order of 10^{26} s. Both of these effects originate from the same supersymmetric operator.

We consider a setup where the fermionic WIMP and its scalar superpartner freeze out with comparable relic densities. While the WIMP itself will have a lifetime greater than the age of the Universe, its scalar partner will decay more rapidly into the dark higgsino state, providing the dominant contribution to the GeV-Matter density, as we now show. If we denote the mass of the light and heavy states by m_L and m_H , respectively, then the relic density of the light and heavy states from freeze-out are $\rho_{L(\text{f.o.})} \propto \sigma_{L(\text{ann.})} \propto m_L^2$ and $\rho_{H(\text{f.o.})} \propto \sigma_{H(\text{ann.})} \propto m_H^2$, respectively, so that $\rho_{L(\text{f.o.})}/\rho_{H(\text{f.o.})} \sim m_L^2/m_H^2$ for roughly equal couplings. After one of the heavy states decays into the light state, the light state will inherit the *number* density of the heavy state, $n_L \sim n_H = \rho_{H(\text{f.o.})}/m_H \propto m_H$, so that the relative relic density from decay $\rho_L/\rho_H \sim m_L/m_H \gg m_L^2/m_H^2$ dominates over $\rho_{L(\text{f.o.})}$. Moreover, for $m_L \sim 5$ GeV and $m_H \sim 1\text{--}2.5$ TeV, the stable heavy state can dominate the dark matter relic density so $\rho_L/\rho_{\text{DM}} \sim m_L/m_H \sim (2\text{--}10) \times 10^{-3}$, which is in precise accord with the DAMA and CoGeNT rates (see Eq. (2)) if the light state has a direct detection cross-section given by (10).

For a concrete example of this scenario, we again consider the low-energy superpotential

$$\mathcal{W} = -\frac{1}{4}W_D^2 - \frac{\epsilon}{2}W_Y^\alpha W_{\alpha D} + \lambda S h_u h_d + M\Phi_u \Phi_d + \frac{1}{2M_1}h_d^2 h_u^2 + \frac{1}{2M_2}h_d^2 \Phi_u^2 \quad (26)$$

where the S , h_u , and h_d are GeV-scale states and the fermions $\tilde{\Phi}_{u,d}$ are the TeV-scale WIMPs. The two higher-dimensional operators induce the necessary splittings among the WIMPs and GeV-Matter, so that in direct detection experiments the GeV-Matter can down-scatter and the heavy dark matter component is forbidden from scattering. Now let us include the Kähler potential terms

$$\int d^4\theta \frac{1}{M_*^2} \left(\Phi_u^\dagger e^V h_u X^\dagger Y + \Phi_u^\dagger e^V h_u X Y^\dagger + \text{h.c.} \right), \quad (27)$$

where X and Y are heavy fields in the supersymmetry breaking sector, and V is the $U(1)_D$ vector field. When X and Y acquire non-zero F_X and F_Y -terms, we break two \mathbb{Z}_2 symmetries under which both superfields Φ_u and Φ_d and either X or Y are charged, allowing the Φ_u scalar to decay through mass mixing with the h_u state. Any non-zero B_μ -term, generated, for example, from gauge mediation, will maximally mix the Φ_u and Φ_d states, and also allow the Φ_d scalar to decay. Two residual (non-supersymmetric) accidental \mathbb{Z}_2 's under which the *fermions* $\tilde{\Phi}_u$ and $\tilde{\Phi}_d$ and either the X or the Y *scalars* are charged, are only broken by the scalar vevs $\langle X \rangle$ and $\langle Y \rangle$, and induces $\tilde{\Phi}_{u,d}$ fermion decay through kinetic mixing with h_u . The fermions are parametrically longer-lived than the scalars, because the higher mass dimension of F_X, F_Y must be compensated for by the TeV-scale Φ mass.

The dominant Φ scalar decay channels are then $\Phi \rightarrow h_u + A'$, $\Phi \rightarrow \tilde{h}_u + \tilde{\lambda}$, and especially $\Phi \rightarrow \tilde{h}_d + \tilde{S}$ due to the fact that $\lambda > g_D$; we estimate the lifetime as

$$\begin{aligned} \tau_\Phi &\sim \left(\frac{4\pi}{\alpha_\lambda} \right) \frac{M_*^4 M^3}{|F_X F_Y|^2} \\ &\sim 10^4 \text{ s} \left(\frac{1/30}{\alpha_\lambda} \right) \left(\frac{M_*}{10^{19} \text{ GeV}} \right)^4 \left(\frac{3 \times 10^7 \text{ GeV}}{\sqrt{F}} \right)^8 \left(\frac{M}{2 \text{ TeV}} \right)^3, \end{aligned} \quad (28)$$

where we have set $F_X = F_Y = F$. The dominant $\tilde{\Phi}$ fermion decays are $\tilde{\Phi} \rightarrow \tilde{h}_u + A'$, $\tilde{\Phi} \rightarrow h_u + \tilde{\lambda}$, $\tilde{\Phi} \rightarrow \tilde{h}_d^\dagger + S^*$, and $\tilde{\Phi} \rightarrow h_d^* + \tilde{S}^\dagger$ with lifetime

$$\begin{aligned} \tau_{\tilde{\Phi}} &\sim \left(\frac{1/30}{\alpha_\lambda} \right) \frac{M_*^4}{|\langle X \rangle \langle Y \rangle|^2 M} \\ &\sim 10^{26} \text{ s} \left(\frac{1/30}{\alpha_\lambda} \right) \left(\frac{M_*}{10^{19} \text{ GeV}} \right)^4 \left(\frac{(2 \times 10^6 \text{ GeV})^2}{\langle X \rangle \langle Y \rangle} \right)^2 \left(\frac{2 \text{ TeV}}{M} \right). \end{aligned} \quad (29)$$

It is important to this argument that less suppressed operators involving X and/or Y and Kähler terms involving chiral products such as $\Phi_u h_d$ are forbidden (as in the UV completion in Appendix A) *or* that one of the scalar vevs $\langle X \rangle, \langle Y \rangle$ vanishes. In either of these situations, the interactions in (27) are the leading decays. In general, all that is required to populate the higgsino states is for scalar Φ decay to occur, so fermion $\tilde{\Phi}$ decays can be rendered negligible if $\langle X \rangle \ll \sqrt{F}$ or $\langle Y \rangle \ll \sqrt{F}$.

The scalar Φ decays occur well after the dark higgs and gauge sectors freeze out, so that the abundance of the GeV-Matter and potentially long-lived (depending on the gravitino mass) $\tilde{\lambda}, \tilde{h}_d$ are determined by the density of Φ . Though initially energetic, the Φ decay products redshift and become non-relativistic before matter-radiation equality. Consequently, the velocity profile of these decay products in galaxy halos is controlled by gravitational infall

like that of the WIMP dark matter. A lifetime in the range $10 \text{ s} \lesssim \tau_\Phi \lesssim 10^5 \text{ s}$ conservatively avoids BBN constraints for $\Omega_\Phi h^2 \simeq 0.1$ and assuming that Φ decays never produce baryons; if the Φ branching fraction to Standard Model leptons is small ($\mathcal{O}(10^{-2})$ is readily achieved from ratios of couplings) then shorter and longer lifetimes are allowed [55].

The fermion decays $\tilde{\Phi} \rightarrow \tilde{h}_u + A'$, $\tilde{\Phi} \rightarrow h_u + \tilde{\lambda}$ and $\tilde{\Phi} \rightarrow h_d^* + \tilde{S}^\dagger$ produce leptons through the decay $A' \rightarrow \ell^+ \ell^-$, $h_d \rightarrow \ell^+ \ell^-$, and $h_u \rightarrow S + A' \rightarrow S + \ell^+ \ell^-$ or $h_u \rightarrow S + \ell^+ \ell^-$ (the latter decay is through an off-shell A' if $m_{h_u} - m_s < m_{A'}$). Consequently, we expect that this model can explain the PAMELA and FERMI signals as arising from $\tilde{\Phi}$ decay for $\langle X \rangle \sim \langle Y \rangle \sim \sqrt{F_X} \sim \sqrt{F_Y} \sim 10^3 - 10^4 \text{ TeV}$, with a possible higher-energy component of the cosmic ray spectrum arising from Sommerfeld-enhanced annihilations (which by itself may not suffice to explain the signal [60]). Note that the $\tilde{\Phi}$ decays also produce high-energy gauginos, which can decay into high energy gravitinos and photons, $\tilde{\lambda} \rightarrow \gamma + \tilde{G}$. If this decay is fast, then it may generate a photon signals close to current limits (see e.g. [61–63]). This constraint is trivially avoided if the gravitino is heavy and/or $\tilde{\lambda}$ is long-lived. (Note that if the gaugino is unstable, its lifetime should also be $\lesssim 10^5 \text{ s}$ to avoid constraints from photodissociation of nuclei [55, 64]) That dark matter decay induced by higher-dimensional operators can produce sizable cosmic ray signals was first pointed out in [65], and more recently considered in [66, 67].

We present a UV completion of the above model (26) and (27) in Appendix A.

V. DIRECT DETECTION PHENOMENOLOGY

In this section, we discuss direct detection signals from light dark matter species at DAMA and CoGeNT, and the constraints from XENON-10 and CDMS-Si. Our focus is on extending previous analyses [26, 27] to models of light dark matter with an inelastic splitting $\delta \lesssim 35 \text{ keV}$ where down-scattering dominates, which (i) changes the recoil and modulation spectra, especially at low recoil energies, (ii) allows lower dark matter masses (down to 3–4 GeV), (iii) permits excellent simultaneous fits to DAMA and CoGeNT spectra and rates, and (iv) greatly reduces the strength of limits from XENON10, and also ameliorates limits from CDMS Si. Down-scattering at lower dark matter masses has also been proposed as an explanation of the DAMA signal [29, 68], and [29] analyzes the phenomenology of a variety of down-scattering signals in detail.

A. Kinematics of Inelastic Down-Scattering

Vector-current interactions of fermions with a small mass splitting are always inelastic. When one component of a Dirac fermion $\Psi = (\psi_L \bar{\psi}_R)$ of mass m receives a small Majorana mass δ , it splits Ψ into two components $\psi_{1,2}$ with Majorana masses $\approx m \pm \delta/2$, which are approximately $\psi_1 = (\psi_L + \psi_R)/\sqrt{2}$ and $\psi_2 = i(\psi_L - \psi_R)/\sqrt{2}$. Upon rotation of the Dirac fermion's gauge interactions to the mass eigenstate basis, all couplings $g_{ij} A^\mu \bar{\psi}_i \bar{\sigma}_\mu \psi_j$ are generated. If the initial state's coupling is vector-like, the coefficients g_{ii} are suppressed by δ/m (as expected because these couplings break the $U(1)_D$ gauge symmetry, and δ parametrizes the $U(1)_D$ breaking felt by the fermion), but if ψ_L and ψ_R have different chiral couplings, the g_{ii} are generically non-zero. An additional suppression of elastic scattering exists *irrespective of the size of g_{ii}* , since the A' couplings to matter are very nearly pure vector-like (with axial couplings down by $(m_{A'}/m_Z)^2$ and proportional to the nuclear spin rather than its total charge), and since the matrix element for a Majorana fermion scattering off a vector current vanishes at zero velocity. Thus, even for the chiral case, *elastic scattering cross-sections are always suppressed by v^2* .

Every halo particle in the high-mass state can down-scatter (unlike up-scattering, there is no threshold kinetic energy). However, the recoil kinematics of down-scattering differs markedly from that of elastic scattering. In the following, we briefly describe the kinematics of down-scattering events for a dark matter particle of mass M_χ scattering off a nucleus of mass m_N . If the dark matter is incident at velocity v , the center-of-mass frame has a velocity $v_{\text{cm}} = \mu v/m_N$ relative to the lab frame, where μ is the dark matter-nucleus reduced mass. The squared velocity of the recoiling nucleus in the center-of-mass frame, assuming the dark matter mass increases by δ in scattering (i.e. $\delta < 0$ for down-scattering) is simply

$$v_{\text{out,cm}}^2 = \frac{\mu^2}{m_N^2} \left(v^2 - \frac{2\delta}{\mu} \right). \quad (30)$$

Therefore, the average lab-frame recoil energy is

$$\bar{E}(v) = \frac{1}{2} m_N \langle v_{\text{lab}}^2 \rangle = \frac{1}{2} m_N (v_{\text{cm}}^2 + v_{\text{out,cm}}^2) = \frac{\mu^2}{m_N} \left(v^2 + \frac{\delta}{\mu} \right), \quad (31)$$

while the range of recoil energies is given by

$$\frac{\mu^2}{2m_N} \left(v - \sqrt{v^2 - \frac{2\delta}{\mu}} \right)^2 \leq E_R \leq \frac{\mu^2}{2m_N} \left(v + \sqrt{v^2 - \frac{2\delta}{\mu}} \right)^2 \quad (32)$$

Thus, when $|\delta| \gtrsim \mu v^2$, the effect of down-scattering on the nuclear recoil spectrum is dramatic, with mean nuclear recoil energy

$$\langle E_R \rangle \approx \frac{|\delta|\mu}{m_N} \quad (33)$$

and width $\Delta E_R \sim (\mu/m_N)\sqrt{8\delta\mu v^2}$.

Because the recoil spectrum is not peaked at zero, experiments with low energy thresholds relative to $|\delta|\mu/m_N$ are sensitive to a large fraction of scattering events, even if the scattering WIMP is quite light (an elastically scattering state of roughly the same mass, such as a dark higgs, can have a comparable halo abundance but go un-noticed because its detection efficiency is typically much lower). On the other hand, experiments with high energy thresholds, which are only sensitive to the tail of the recoil energy distributions, are less sensitive to down-scattering than to elastic scattering with the same characteristic scale for E_R . This is because, in the elastic case, the typical E_R and its exponential tail are both determined by $(\mu^2/m_N)v^2$. Down-scattering models achieve the same characteristic E_R scale with lower dark matter masses, and therefore a sharper exponential fall-off in rate at high recoil energies. This greatly reduces the sensitivity of Xenon searches and slightly ameliorates the constraints on light dark matter from CDMS Silicon detectors. Moreover, the bounds from these experiments are extremely sensitive to uncertainties in the threshold energy and the detection efficiency near threshold.

B. Direct Detection Scattering Rate

The differential rate for spin-independent scattering of a dark matter species χ off a target containing N_T nuclei of mass m_N is given by

$$\begin{aligned} \frac{dR}{dE_R} = & N_T \frac{m_N f_\chi \rho_{\text{dm}} (f_p Z + f_n (A - Z))^2}{2\mu_n^2 m_\chi f_p^2} \sigma_{\chi,p} \\ & \times F^2(E_R) \int_{v_{\text{min}}}^{v_{\text{esc}}} d^3v \frac{f(\vec{v} - \vec{v}_e)}{v}, \end{aligned} \quad (34)$$

where E_R is the energy of the recoiling nucleus, μ_n is the dark matter-nucleon (proton or neutron) reduced mass, $\rho_{\text{dm}} \simeq 0.3 \text{ GeV}/\text{cm}^3$ is the local dark matter density, m_χ and f_χ are the mass and fractional abundance $\rho_\chi/\rho_{\text{dm}}$ of χ , Z and A are the atomic number and atomic mass of the target nucleus, f_p and f_n are the effective coupling strengths of the dark matter

particle to protons and neutrons, respectively (for kinetic mixing, $f_n = 0$ so the scattering rate is proportional to Z^2), and

$$\sigma_{\chi,p} \equiv \frac{\mu_n^2}{64\pi^2 m_{DM}^2 m_p^2} \int d\Omega_{cm} |\mathcal{M}|^2 \quad (35)$$

is the effective scattering cross-section of the dark matter particle off protons (this equals the physical scattering cross-section in the elastic limit, but differs for inelastic scattering by a ratio of outgoing to incoming velocities in the CM frame). The parametrization of the form factor $F(E_R)$, dark matter velocity distribution $f(\vec{v} - \vec{v}_r)$, and E_R -dependent velocity threshold v_{min} in the last two terms of (34) are described in Appendix C.

Throughout this section, we treat $\sigma_{\chi,p} f_\chi$ as the parameter to be fit or constrained by event rates in direct detection experiments. For the inelastic scenarios, f_χ refers to the *total* χ mass fraction, which we assume to be equally divided between the excited and ground states, so that the fraction of the halo that can down-scatter is $f_\chi/2$, while the fraction that can up-scatter is $f_\chi/2$ times a Boltzmann factor, and we add the two contributions.

C. The Data

CoGeNT CoGeNT is an ultra-low-noise germanium detector with a fiducial mass of 330 g that operated for 56 days in the Soudan Underground Laboratory with an ionization energy threshold of 0.4 keVee (keV electron equivalent, i.e. ionization energy). Between 0.4 keVee and 3.2 keVee, the detector recorded an approximately uniform background rate, sharp cosmogenic peaks, and about 100 events of unknown origin between 0.4 and 1.0 keVee, peaked near threshold [26]. Background explanations of the latter component are not excluded, but its spectrum is sufficiently similar to that expected from light dark matter scattering off germanium to warrant further study [26–28, 32].

The proximity of the putative signal to the detector’s noise threshold and the similarity of its shape to that of a hypothetical exponential background pose a challenge for interpretation of the CoGeNT dataset. We do not fit to the first bin (0.4-0.45 keVee) on the grounds that it is within 1σ energy resolution from the noise threshold and may be significantly contaminated. We fit three contributions to the remaining bins between 0.45 and 3.2 keVee:

1. A uniform background with free normalization.

2. Gaussian peaks centered at 1.1 keVee and 1.29 keVee, originating from known radioactive decays, each of which is modeled as a Gaussian with free normalization and width given by the detector energy resolution $\delta E \simeq \sqrt{(0.0697)^2 + 0.000977 E_R}$ keV from [69].
3. A contribution from dark matter scattering calculated according to (34), using the time-average earth velocity. We assume the ionization energy (in keVee) from a nuclear recoil of energy E_R is given by

$$E_R = E_{ee}/\text{QF} \quad (36)$$

with quenching factor $\text{QF}(\text{Ge})=0.29$. We do not include any possible channeling in Ge. We neglect the finite energy resolution as it is much narrower than the intrinsic energy distribution of dark matter recoils.

All expectations are multiplied by the detection efficiency plotted in [69] as a function of energy. As the number of events in each bin is small, we use Poisson log-likelihood rather than χ^2 to quantify goodness of fit. It is important to remember that the CoGeNT low-energy excess is entirely consistent with an exponentially falling background. However, we do not find it useful to include such a component when fitting the spectrum because it is not distinguishable from a dark matter signal with the present experimental statistics. At higher statistics, a spectral feature (which is expected in much of the down-scattering parameter space) could help to establish the dark matter hypothesis if it is correct.

DAMA/NaI and DAMA/LIBRA The DAMA/NaI [23] and DAMA/LIBRA [24, 25] experiments have reported an annual modulation in the rate of single-hit events with ionization energy in the interval 2–5 keVee, with a consistent phase over 13 years. Here we fit this signal assuming it originates from dark matter scattering off the Sodium and Iodine nuclei in their detectors.

The detector response to nuclear recoils is complicated by the “channeling effect” [70], whereby a fraction of recoiling nuclei deposit almost all of their energy electronically. Following the parametrization of [71], we take $E_{ee} = E_R$ in a fraction f_{chan} of scattering events (the “channeled events”), and (36) in the remainder, with

$$f_{\text{chan,Na}} \simeq \frac{e^{-E_R/18}}{1 + 0.75E_R}, \quad f_{\text{chan,I}} \simeq \frac{e^{-E_R/40}}{1 + 0.65E_R}, \quad \text{QF}(\text{Na}) = 0.3, \quad \text{and} \quad \text{QF}(\text{I}) = 0.09. \quad (37)$$

Thus, some fraction of events at 2–5 keVee arise from 2–5 keV nuclear recoils. In the models we consider, channeled I recoils dominate the expected signal, with a smaller contribution from un-channeled Na recoils.

It is important to note that the functional forms in (37) are parametrizations of the Monte Carlo results in [49], which differ significantly from analytical models [72, 73] and have not been tested experimentally in the relevant energy range. Moreover, partial channeling is not included in these parametrizations. All of our quantitative predictions for the DAMA experiment depend very sensitively on the behavior of channeling in the 2-5 keV range.

We calculate the modulation spectrum by subtracting the expected energy spectrum when the Earth’s velocity in the galactic frame is at its annual minimum, from the expected spectrum when this velocity is at its annual maximum. We smear this spectrum with a gaussian whose width corresponds to the DAMA energy resolution, $\delta E = 0.0091 + 0.448/\sqrt{E}$ [74].

XENON10 The XENON10 experiment used a liquid Xenon target with an exposure of 316 kg-days, and saw no events consistent with light dark matter. We use the 2–20 keV nuclear recoil energy bin and the reported software cut acceptances and nuclear recoil band acceptance in [44] to set a limit.

A large source of uncertainty comes from the scintillation efficiency, L_{eff} , which corresponds to the amount of nuclear recoil energy that gets recorded as scintillation light. This affects the detector’s threshold energy, and a small change in L_{eff} has a large impact on the resulting constraints. Unless otherwise stated, we use the recent measurement of L_{eff} in [45], which yields weaker constraints than the value $L_{\text{eff}} = 0.19$ quoted in [44].

CDMS-Si The CDMS experiment contains both Germanium and Silicon detectors. Far less data has been collected for the Silicon detectors, but we have checked that this data nonetheless sets a tighter constraint than the recent Germanium analysis [75]. For CDMS-Si, a total exposure of 53.5 kg-days has been obtained with a threshold of 7 keV nuclear recoil [47] (we use this updated result rather than the results in [76]). We use the efficiency provided in [47]. While the resulting constraints generate some tension with the CoGeNT and DAMA preferred regions, they are extremely sensitive to the precise threshold energy, as we will demonstrate.

D. Results

Figure 1 shows the regions consistent with the CoGeNT and DAMA data for different choices of the mass splitting δ from 0 to 30 keV. We take $f_p = 1$ and $f_n = 0$, as appropriate for dark matter interacting with nuclei through a kinetically mixed mediator. The elastic scattering regions (green region and contours) are consistent with the results for $f_p = f_n = 1$ in [26, 27], after rescaling our results by a factor of $(Z(\text{Ge})/A(\text{Ge}))^2 \simeq 0.19$. The CoGeNT regions correspond to regions where $-2\Delta \ln \mathcal{L} + 2 \ln \mathcal{L}_{\max} = 6.25$ (and 14.16 for the right plot), where \mathcal{L} is the Poisson likelihood and \mathcal{L} is its maximum value for *any* $|\delta|$ (here near $\delta = 0$). Likewise the DAMA contours correspond to $\chi^2 - \chi_{\min}^2 = 6.25$ (14.16) relative to the overall minimum χ^2 , occurring near $|\delta| = 20$ keV. These boundaries approximate formal 90% (1.6σ) and 99.7% (3σ) likelihood regions, but they neglect systematic uncertainties in the dark matter halo and detector response that are known to be important.

We note that the down-scattering parameter space allows simultaneous fits to the CoGeNT and DAMA data. In our parametrization of channeling, the preferred region is

$$m_{\text{dm}} \sim 4 - 8 \text{ GeV}, \quad \sigma_0 \times f \sim (1 - 2) \times 10^{-40} \text{ cm}^2, \quad |\delta| \sim 15 - 25 \text{ keV}. \quad (38)$$

but would shift according to any change in the treatment of channeling.

Figure 2 shows expected spectra at CoGeNT and DAMA for two benchmark points. The top two panels show the recoil spectra for $|\delta| = 20$ keV, $m_{\text{dm}} = 6.0$ GeV, and $\sigma \cdot f \simeq 1.4 \times 10^{-40} \text{ cm}^2$, a point in this preferred region (indicated by the dot in Figure 1). The CoGeNT recoil spectrum agrees with (33) after accounting for the quenching factor $\text{QF}(\text{Ge})=0.29$. The modulation at DAMA comes from the high-energy tail of the nuclear recoil energy distribution, and is dominated by channeled iodine recoils (with a smaller contribution from unchanneled sodium recoils). The large component below the 2–5 keV DAMA region of interest (and below the 2 keV threshold of the DAMA modulation analysis) comes from unchanneled iodine scatters. For these parameters, we expect 6 events in CDMS-Si and 2 in XENON10 under our baseline assumptions. Under different assumptions about the DAMA detector response, the preferred regions move (and can reduce the tension with CDMS-Si and XENON10). The lower panels show a point with $|\delta| = 30$ keV, $m_{\text{dm}} = 4.5$ GeV, and $\sigma \cdot f \simeq 0.95 \times 10^{-40} \text{ cm}^2$, which provides a good fit to CoGeNT and for which we expect only 1.6 events in CDMS-Si and none at XENON10. Although the fit to the DAMA recoil spectrum with baseline parameters (red curve in lower right panel of Figure 2) is worse than

before, it can be remedied by, for example, increasing the channeled Iodine fraction by 20% and by shifting the overall DAMA energy scale by 10%. The resulting dashed blue curve fits very well.

Constraints CDMS-Si, XENON10, and the DAMA unmodulated single-hit spectrum all have the potential to constrain these models. However, significant systematic uncertainties apply to each of these limits. Curves corresponding to 2.3 and 5 expected events in XENON10 and CDMS-Si data, with the baseline assumptions detailed above, are included in the right panel of Figure 1. These limits, especially the CDMS-Si limit, seem to be in some tension with the preferred overlapping region of DAMA and CoGeNT. However, their severity is rather sensitive to various assumptions about each detector’s response near threshold and even the threshold energy, while the precise regions preferred by CoGeNT and DAMA are vulnerable to similar uncertainties (and particularly, in the case of DAMA, to the detailed features of channeling). We illustrate the sensitivity of the XENON10 and CDMS-Si limits to the energy scales of these experiments in Figure 3, where we have varied the true energy scale of their detector thresholds by 20%.

The potential importance of the DAMA unmodulated single-hit rate [24] (which varies slowly between 1 and 1.5 counts/kg/day for energies above 1.5 keVee and rises rapidly at lower energies) in constraining models of light dark matter was emphasized in [43]. With the standard channeling assumptions, our benchmark points overpredict the total rate near 2 keVee by about a factor of 2. This is a potentially severe source of tension, but it is driven by two poorly constrained features of channeling — the behavior of channeling at low ($\lesssim 3$ keVee) energies, and a possible energy dependence of the quenching factor in this range. Blocking and de-channeling effects can dramatically reduce the channeling fraction at low energies, rendering the predicted single-hit rate consistent [72]. A careful study of channeling in the few-keV region is needed to make sharp predictions of these rates and spectral profiles in order to settle this issue. The single-hit tension is also driven by fitting the *dip* in the lowest DAMA bin. Allowing for a mild uncorrected decrease in efficiency in this bin (so that the rate in this bin is comparable to the 2nd bin) would render lower splittings consistent with the data, and make the unmodulated spectrum consistent. We also note that the lower-mass parameter regions discussed in [29], which must assume Ge channeling to explain the CoGeNT excess, do not have this problem.

Finally, although we have checked that the preferred regions and constraints depend only

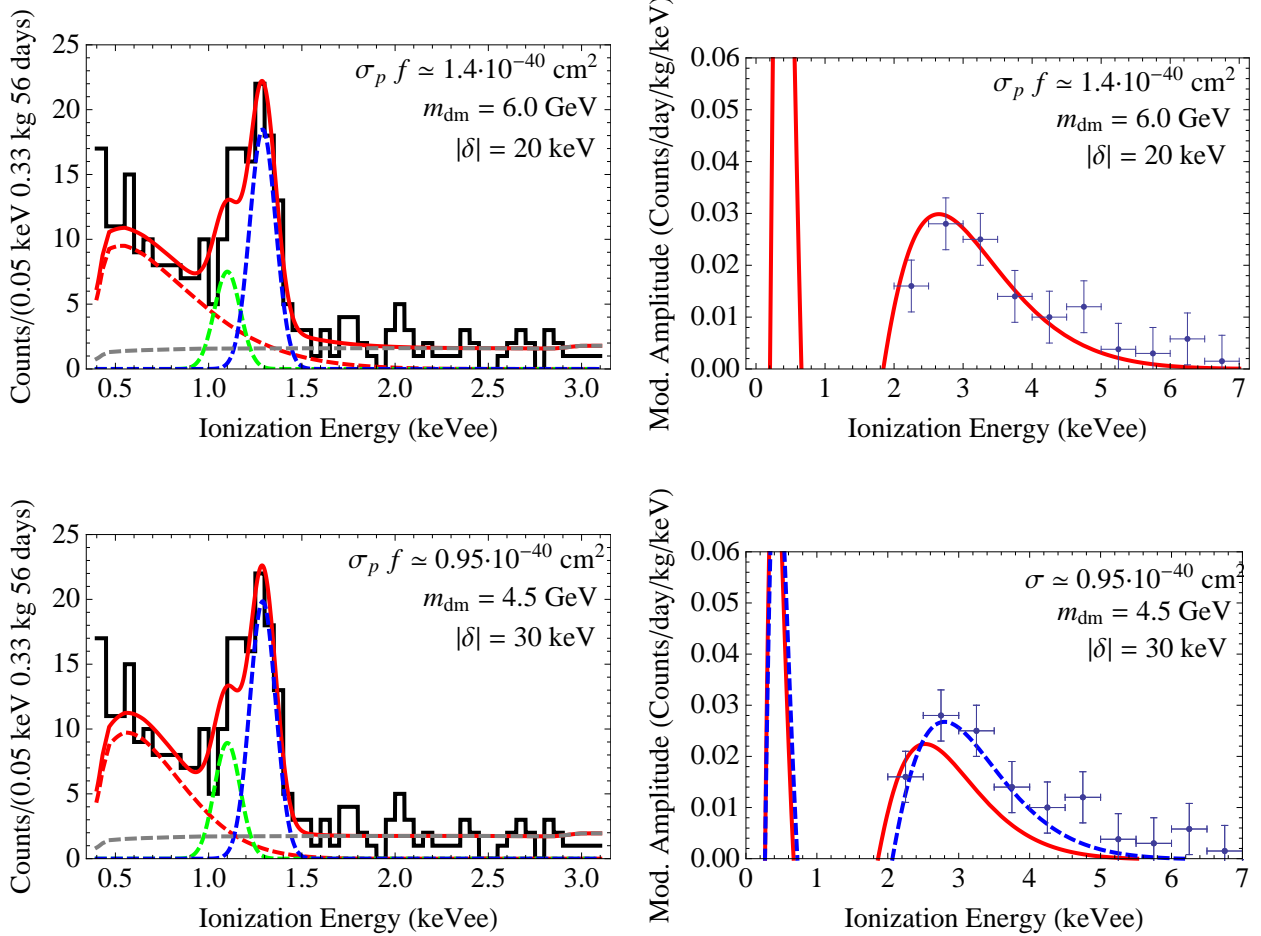


FIG. 2: **Top:** A fit (solid red line) to the CoGeNT data (left) and DAMA data (right). In the CoGeNT fit, we assume an energy-dependent detection efficiency [26] and include four contributions: dark matter (red dashed), two gaussians for the cosmogenic peaks near 1.1 keV and 1.29 keV (green and blue dashed), and a constant background (gray dashed). The dark matter mass (5.5 GeV), cross-section ($1.4 \times 10^{-40} \text{ cm}^2$), and splitting $\delta = 20 \text{ keV}$ correspond to a benchmark point that fits both data sets very well (indicated by the dot in Figure 1). The expected number of events in CDMS-Si and XENON10 is about 6 and 2, respectively. Assuming a 20% larger energy threshold in either experiment would lead to 2 and 0 events expected, respectively. **Bottom:** The same as the top two figures but now for a dark matter mass of 4.5 GeV, a cross-section of $0.95 \times 10^{-40} \text{ cm}^2$, and a splitting of $\delta = 30 \text{ keV}$ (upside-down triangle in Figure 1). This leads to ~ 1.6 events in CDMS-Si and 0 in XENON10. The agreement with DAMA data (solid red curve in right figure) is improved by assuming an increase of 20% in the Iodine channeling fraction and a shift of the overall DAMA energy scale by 10%, producing the blue dashed curve.

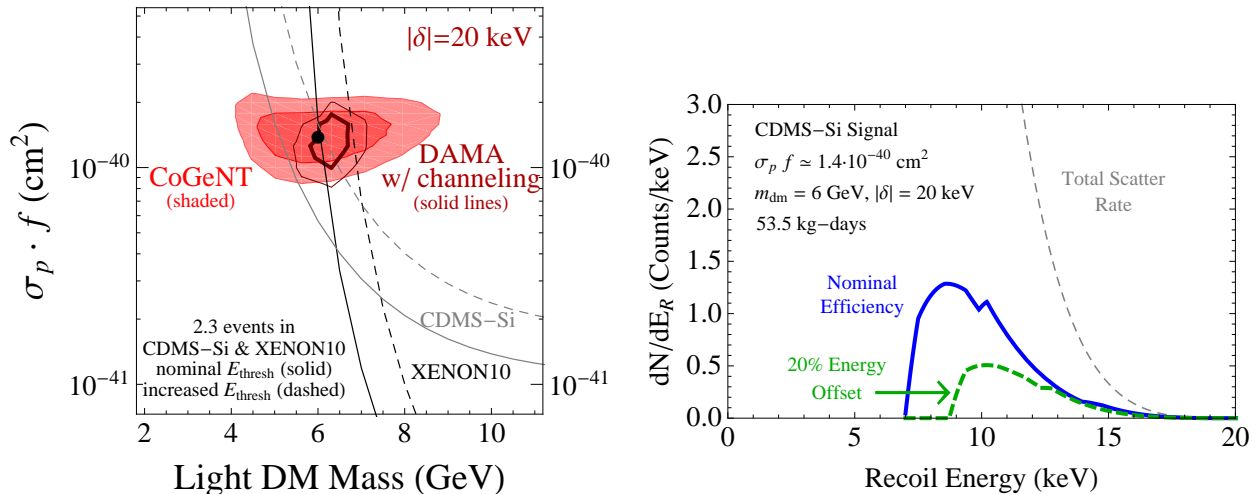


FIG. 3: The effect of energy thresholds on XENON10 and CDMS-Si Limits. **Left:** The contours and thin solid lines are as in the left panel of Figure 1. The dashed black and gray lines indicate the change in XENON10 and CDMS-Si sensitivity when their threshold energy is increased by 20%. In the case of CDMS-Si, we use the same rescaling of recoil energies in determining the detection efficiency. **Right:** The spectra in CDMS-Si for the benchmark point with $\delta = 20$ keV (top panels of Figure 2), assuming the nominal efficiency (blue solid curve) or assuming a 20% offset of the energy threshold and energy-dependent efficiency (dashed green curve). The total scatter rate with no efficiency correction applied (thin dashed gray curve) is also shown for reference.

mildly on the mean halo velocity, and are quite insensitive to variations in v_{esc} , other changes to the halo velocity profile may have a more significant impact.

VI. B-FACTORY PREDICTIONS

For all of the scenarios outlined above, the mass of the vector A' and higgs states are underneath the energy threshold for direct production at the B -factories BaBar and Belle [34, 35]. The main production processes are $e^+e^- \rightarrow A'\gamma$ and $e^+e^- \rightarrow A'h_d$. For $m_{A'} \sim 1$ GeV, the dominant decay modes of A' are e^+e^- , $\mu^+\mu^-$ or $\pi^+\pi^-$, all of which can have branching fractions $\gtrsim 10\%$. Other decays with less than $\lesssim 10\%$ branching fractions include, $\pi^0\omega$, K^+K^- , $\pi^+\pi^-2\pi^0$, $2\pi^+2\pi^-$, $\pi^0\pi^+\pi^-$, and $K^0\bar{K}^0$. The h_d will mainly decay to an A' and two soft leptons or pions, or sometimes directly to two muons or pions [34, 35]. Other hadronic final states are also possible for $m_{h_d} \gtrsim 1$ GeV. From [34, 35], the production cross-

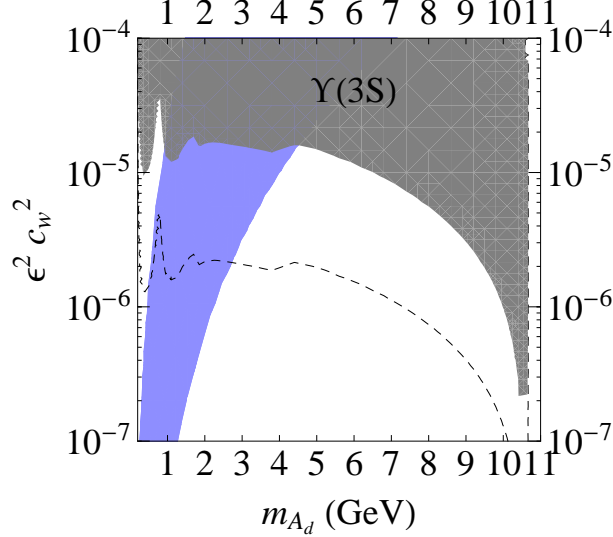


FIG. 4: Constraints derived from the $e^+e^- \rightarrow A'\gamma$ process in the $\mu^+\mu^-\gamma$ search in BaBar's $\Upsilon(3S)$ data using [40]. The gray region is excluded at more than 90% confidence, while the dotted line shows the combined reach of BaBar and Belle. The blue region shows the expected range of parameters in models of light sub-dominant higgsino dark matter. While not shown, we expect that multi-lepton searches sensitive to the reaction $e^+e^- \rightarrow A'h_d$ in the four or more lepton/pion final state can reach sensitivities of order $(\epsilon c_W)^2 \sim \mathcal{O}(10^{-8})$ [77].

sections are,

$$\sigma_{e^+e^- \rightarrow A'\gamma} \approx 60 \text{ fb} \left(\frac{\epsilon^2 c_W^2}{10^{-5}} \right) \left(\frac{E_{cm}^2 - m_{A'}^2}{E_{cm}^2} \right)^{-1}, \quad (39)$$

$$\sigma_{e^+e^- \rightarrow A'h_d} \approx 2 \text{ fb} \times \left(\frac{\alpha_d}{\alpha} \right) \left(\frac{\epsilon^2 c_W^2}{10^{-5}} \right) \left(\frac{E_{cm}^2 - (m_{A'} + m_{h_d})^2}{E_{cm}^2} \right)^{-1}, \quad (40)$$

valid over the range $1 \text{ GeV} \lesssim m_{A'} + m_{h_d} < E_{cm}$. See [34, 35] for a detailed discussion of these reactions.

Assuming that the branching fraction for the decay $A' \rightarrow \mu^+\mu^-$ is $1/(1 + R_{\mu^+\mu^-}) \lesssim 30\%$ for $m_{A'} \gtrsim 1 \text{ GeV}$, current limits on $e^+e^- \rightarrow A'\gamma \rightarrow \mu^+\mu^-\gamma$ require $(\epsilon c_W)^2 \lesssim 2 \times 10^{-5}$ over the range $1 \text{ GeV} \lesssim m_{A'} \lesssim 6 \text{ GeV}$ [40], and are somewhat stronger at higher mass. Figure 4 illustrates the limit on $(\epsilon c_W)^2$ that we derived from [40]. Limits on $e^+e^- \rightarrow 4l$ [77] and $e^+e^- \rightarrow \text{invisible} + \gamma$ [78, 79] can also be constraining, but are slightly weaker than the di-lepton channel for the range of parameters we are considering. Likewise, rare decay constraints are also not as strong as the di-lepton search [38, 80].

In the models considered in this paper, $(\epsilon c_W)^2$ is predicted as a function of α_D and $m_{A'}$. An improved search for $e^+e^- \rightarrow A'\gamma$ in the $\mu^+\mu^-\gamma$ final state using the full $\Upsilon(4S)$ data set at BaBar should be sufficient to probe almost the entire parameter space of this class of theories. Even better is a combined analysis of BaBar and Belle data in this channel. If kinematically accessible, the $e^+e^- \rightarrow A'h_d$ process can offer a particularly striking signature in the six lepton (and pion) final state. Existing searches in four lepton channels [77] indicate that upcoming searches in higher lepton (and pion) multiplicity channels should thoroughly explore this region.

For the scenario where the relic abundance of higgsino GeV-Matter arises from thermal freeze-out, normalizing $\alpha_D \propto m_{A'}^2/m_{\tilde{h}}$ to the CoGeNT signal implies a particularly sharp prediction for $(\epsilon c_W)^2$, given by (21). The predicted B -factory cross-sections in this case are remarkably close to current limits [40, 77–79] and are readily tested by extending existing B -factory searches.

While we believe that B -factories are uniquely suited to explore the light dark matter scenarios outlined in this paper, LEP and hadronic collider searches for lepton jets can also play an important role. If Standard Model super-partners are light enough to be directly produced at colliders, then supersymmetric dark sectors can give rise to spectacular lepton jet signals at the Tevatron or LHC [81–85]. However, note that the super-partner scale can be as high as $\sim 1/\sqrt{\epsilon}m_W \sim 10$ TeV before irreducible super-symmetry breaking feeds into the dark sectors we consider. Additional searches in J/Ψ -factories [41, 42], ϕ -factories [36, 37], fixed-target experiments [86–93], and satellite searches [61, 94, 95] can provide evidence for (or constrain) GeV-scale gauge sectors coupled through kinetic mixing.

VII. CONCLUSION

If supersymmetry stabilizes the weak scale, then kinetic mixing with a new gauge force can naturally generate a GeV-scale sector. This remarkable possibility could open the door to exploring supersymmetry at low energies. In simple supersymmetric dark sectors, stable GeV-Matter can arise in the dark higgs sector, and these may comprise a portion of the dark matter with new and interesting phenomenology. This possibility is well-motivated by cosmic ray data suggestive of GeV-scale forces.

In this paper, we have proposed simple models of GeV-Matter within a supersymmetric

dark sector that couples to the Standard Model through kinetic mixing. The low-energy effective Lagrangian we have studied has already appeared in the literature to realize inelastic up-scattering of $\mathcal{O}(100 \text{ GeV})$ mass dark matter as an explanation of the DAMA signal. Our interpretation of the data is different:

- The CoGeNT and DAMA direct detection anomalies are simultaneously explained by inelastic *down-scattering* of $\mathcal{O}(\text{GeV})$ mass dark-sector states (GeV-Matter), for which the higgsinos of a supersymmetric dark sector are an excellent candidate. These states comprise a very sub-dominant fraction of the dark matter. Tension with existing experiments is reduced, and the CoGeNT and DAMA rates readily agree — in contrast to elastic scattering explanations.
- Electroweak symmetry breaking triggers breaking of the dark sector $U(1)_D$, giving the vector multiplet a mass $m_{A'} \sim \sqrt{\epsilon g_D} m_W \sim \text{GeV}$. The direct detection cross-section of GeV-Matter is predicted (equation (10)) in terms of Standard Model parameters with no dependence on ϵ or α_D .
- The correct CoGeNT and DAMA event rates are obtained for GeV-Matter comprising 0.2–1% of the halo mass density. This abundance can be explained by thermal freeze-out, and the couplings required for this scenario are readily tested with existing B -factory data.
- Another possibility is that a metastable WIMP decays to GeV-Matter after thermal GeV-Matter annihilation has frozen out. In this case, $\frac{\rho_{\text{light}}}{\rho_{\text{WIMP}}} \approx \frac{m_{\text{light}}}{m_{\text{WIMP}}}$ robustly produces the desired relic abundance for metastable WIMP masses of 2–3 TeV. We have exhibited a model where the metastable WIMP is the bosonic superpartner of the TeV-mass dark matter whose decays can explain the cosmic ray data.

Thus, in our unified and predictive framework, the GeV-scale dark sector plays a crucial role in explaining cosmic ray excesses *and* contains stable GeV-Matter responsible for direct detection signals. The most important test of our interpretation of DAMA and CoGeNT will come from further searches with direct detection experiments. In particular, further studies using CDMS Silicon and Germanium data, especially with lower threshold energy, will be vital. The power of these experiments depends sensitively on their threshold energies and signal efficiencies. Future studies should include these uncertainties when quoting

constraints. Low-energy flavor factories play an equally important role in testing our proposal. A resonance search in $e^+e^- \rightarrow \gamma\mu^+\mu^-$ in the full BaBar or Belle datasets and higher lepton-multiplicity searches should be powerful enough to discover or exclude the light vector A' over most of the parameter space that can explain CoGeNT and DAMA. Additional searches in J/Ψ -factories [41], ϕ -factories [36, 37], and fixed-target experiments [86–92] can provide evidence for (or constrain) a GeV-scale gauge sector coupled through kinetic mixing. Finally, if Standard Model super-partners are light enough to be directly produced at colliders, then supersymmetric dark sectors can give rise to spectacular lepton jet signals at the Tevatron or LHC [82–84, 96].

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Appendix A: A UV completion for GeV-Matter from Decay

For completeness, we present a simple UV completion of the model of Section IV, where the GeV-Matter states are populated from the decay of a TeV scale state. Consider the following matter content:

	S	Φ_u	Φ_d	h_u	h_d	Z_1	Z_2	X	Y	A	\bar{A}
$U(1)_d$	0	1	-1	1	-1	0	0	0	0	-1	1
$U(1)_R$	1	0	2	0	1	1	1	φ	φ	$2 - \varphi$	φ
\mathbb{Z}_2	0	-1	-1	0	0	0	-1	-1	0	0	0

where φ is a real number chosen in such a way to guarantee the form of the superpotential and Kähler potential corrections. At some very high scale, the complete superpotential is given by

$$\mathcal{W} = W_{\text{MSSM}} + W_{\text{dark}} + W_{\text{mix}} + W_{\text{split}} + W_{\text{K}} + W_{\text{break}}, \quad (\text{A1})$$

where W_{MSSM} is the supersymmetric Standard Model,

$$W_{\text{dark}} = -\frac{1}{4}W_D^2 + \lambda S h_u h_d + M \Phi_u \Phi_d \quad (\text{A2})$$

contains the GeV-scale dark sector fields (including the higgsinos, which will form the GeV-Matter) and the TeV-scale dominant dark matter component from the $\Phi_{u,d}$, and

$$W_{\text{mix}} = -\frac{\epsilon}{2}W_Y^\alpha W_{\alpha D} \quad (\text{A3})$$

contains the gauge-kinetic mixing term that connects the supersymmetric Standard Model to the dark sector. In addition,

$$W_{\text{split}} = Z_1 h_u h_d + Z_2 \Phi_u h_d - \frac{1}{2}M_1 Z_1^2 - \frac{1}{2}M_2 Z_2^2 \quad (\text{A4})$$

will, after integrating out the \sim TeV-mass fields Z_1 and Z_2 , generate the higher dimensional terms in Eq. (26) that generate the mass splittings among the heavy and light dark matter states. The superpotential

$$W_K = AX\Phi_u + AYh_u + M_*A\bar{A} \quad (\text{A5})$$

will, after integrating out the fields A and \bar{A} at some high scale M_* , generate the Kähler potential term Eq. (27), but no additional superpotential terms. Note that the symmetries guarantee that the Kähler potential has no lower-dimensional terms, or other terms of the same dimension as Eq. (27). Finally, we will not specify the precise form of W_{break} . This superpotential contains fields that break supersymmetry to give non-zero F-terms F_X and F_Y . It also has to contain the supersymmetry breaking sector that generates the soft masses in the supersymmetric Standard Model via some mediation mechanism.

Appendix B: Sources of WIMP and GeV-Matter Mass Splittings

We briefly remark on three models that give rise to inelastic splittings for the Φ fermions and dark higgsinos. The simplest model has $\Phi_{u,d}$ charges $\pm 1/2$, so that additional interactions

$$\mathcal{W}' = \lambda_+ h_d \Phi_u^2 + \lambda_- h_u \Phi_d^2 \quad (\text{B1})$$

are possible. The former directly generates a somewhat large Majorana mass for the heavy dark matter, while integrating out Φ loops or other heavy states generates an operator of the form (16) and therefore a splitting of the dark higgsinos.

Smaller hierarchies of splittings are achieved in models where $\Phi_{u,d}$ have charge ± 1 . We first consider a model with two singlets X and Y , and two \mathbb{Z}_2 global symmetries: one under which X , S , and h_u are odd and one under which Y , S , and h_d are odd. These symmetries allow, in addition to (5) and (11), relevant and marginal interactions

$$\mathcal{W}_{XY} = \mu_X X^2 + \mu_Y Y^2 + \lambda_- \Phi_d h_u X + \lambda_+ \Phi_u h_d Y + \lambda_s SXY. \quad (\text{B2})$$

Integrating out X and Y gives rise to interactions

$$\frac{1}{M} \Phi_u^2 h_d^2 \quad \text{and} \quad \frac{1}{16\pi^2} \frac{1}{M} h_u^2 h_d^2, \quad (\text{B3})$$

which split the heavy dark matter and dark higgsinos respectively. The Φ splitting is generated at tree-level and the dark higgsino splitting only at loop-level, so the dark higgsino splitting is naturally smaller. For Yukawa couplings $\sim 1/10$, $\langle h_d \rangle \sim 10$ GeV, and $M \sim 1$ TeV the typical splittings are 1 MeV and 10 keV respectively for the heavy and light states.

Alternately, our Φ and $U(1)_D$ higgs-sector fields can couple to a doubly-charged field $T_{\pm 2}$ through interactions

$$\mathcal{W}_T = \mu_T T_{+2} T_{-2} + \lambda_d h_d^2 T_{+2} + \lambda_u h_u^2 T_{-2} + \lambda_- \Phi_d^2 T_{+2} + \lambda_+ \Phi_u^2 T_{-2}. \quad (\text{B4})$$

After $U(1)_D$ symmetry-breaking, the F -flat condition for T_{+2} is satisfied if T_{-2} acquires a small vev $\langle T_{-2} \rangle = -\lambda_d v_d^2 / \mu_T \sim \text{MeV}$ for $\mu_T \sim \text{few TeV}$. This vev induces small Majorana masses for both Φ_d and h_d ; a hierarchy of Yukawa couplings $\lambda_u \ll \lambda_+$ leads to a hierarchy of splittings.

Appendix C: Form Factor and Halo Parametrizations for Direct Detection

We elaborate here on the form factor and halo parametrizations appearing in (34), which are used throughout Section V. $F(E_R)$ is a nuclear form factor, which we model as a Helm form factor as in [98],

$$F(qr_n) = \frac{3j_1(qr_n)}{qr_n} e^{-(qs)^2/2}, \quad (\text{C1})$$

with $q = \sqrt{2m_N E_R}$, $r_n^2 = c^2 + \frac{7}{3}\pi^2 a^2 - 5s^2$, $c = 1.23A^{1/3} - 0.60$ fm, $s = 0.9$ fm, and $a = 0.52$ fm.

In the final factor of (34),

$$\int_{v_{\min}}^{v_{\text{esc}}} d^3v \frac{f(\vec{v} - \vec{v}_e)}{v}, \quad (\text{C2})$$

\vec{v} is the dark matter velocity in the frame of the Earth, \vec{v}_e is the velocity of the Earth in galactic coordinates, and $f(\vec{v} - \vec{v}_e)$ is the local dark matter velocity distribution in the Earth rest frame, which we model by

$$f(\vec{v} - \vec{v}_e) \propto \left(e^{-(\vec{v} - \vec{v}_e)^2/v_0^2} - e^{-v_{esc}^2/v_0^2} \right), \quad (\text{C3})$$

for $v < v_{esc}$ and 0 otherwise. We normalize $f(\vec{v} - \vec{v}_e)$ to unity (for example, if $v_{esc} = \infty$, then the proportionality constant is $1/(\pi v_0^2)^{3/2}$). We use $v_0 = 254$ km/s [99] for both the average velocity of local stars and the mean velocity of halo particles. The local galactic escape velocity v_{esc} is measured to lie in the range 498 km/s $< v_{esc} < 608$ km/s at 90% confidence [100], and we use the median likelihood value of $v_{esc} \simeq 544$ km/s. The Earth velocity in the galactic frame is determined as a function of time by $\vec{v}_e = \vec{v}_\odot + \vec{v}_\oplus$, where \vec{v}_\odot is the Sun's velocity relative to the Galactic rest frame [99, 101],

$$\vec{v}_\odot = \begin{pmatrix} 10.00 \\ 5.23 \\ 7.17 \end{pmatrix} \text{ km/s} + \begin{pmatrix} 0 \\ v_0 \\ 0 \end{pmatrix} \quad (\text{C4})$$

and \vec{v}_\oplus is the Earth's velocity relative to the Sun,

$$\vec{v}_\oplus = \langle u_E \rangle (1 - e \sin(\lambda - \lambda_0)) \begin{pmatrix} \cos(\beta_x) \sin(\lambda - \lambda_x) \\ \cos(\beta_y) \sin(\lambda - \lambda_y) \\ \cos(\beta_z) \sin(\lambda - \lambda_z) \end{pmatrix} \text{ km/s}. \quad (\text{C5})$$

Here, the Earth's orbit has a mean velocity $\langle u_E \rangle = 29.79$ km/s and ellipticity $e = 0.016722$, λ is the angular position of the Earth's orbit, and the quantities $\beta_{x,y,z}$ and $\lambda_{0,x,y,z}$ are given in [98].

The velocity integral in (34) cuts off at a minimum velocity v_{min} determined by the scattering kinematics:

$$v_{min} = \frac{1}{\sqrt{2m_N E_R}} \left| \frac{m_N E_R}{\mu_N} + \delta \right|, \quad (\text{C6})$$

where μ_N is the dark matter-nucleus reduced mass, and δ is the splitting between the incoming and outgoing dark matter particles: $\delta = 0$ for elastic scattering, $\delta > 0$ if the incoming dark matter particle scatters into a heavier excited state (“up-scattering”) [102], while $\delta < 0$ if the incoming dark matter particle is already in an excited state and scatters

into a lighter ground state (“down-scattering”). For the masses and splittings we consider, at most a tiny fraction of the halo can up-scatter and the down-scattering signal dominates.

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- [1] O. Adriani et al. (PAMELA) (2008), 0810.4995.
 - [2] A. A. Abdo et al. (The Fermi LAT) (2009), 0905.0025.
 - [3] F. Aharonian et al. (H.E.S.S.), *Phys. Rev. Lett.* **101**, 261104 (2008), 0811.3894.
 - [4] F. Aharonian (H.E.S.S.) (2009), 0905.0105.
 - [5] D. P. Finkbeiner, *Astrophys. J.* **614**, 186 (2004), astro-ph/0311547.
 - [6] D. P. Finkbeiner (2004), astro-ph/0409027.
 - [7] D. Hooper, D. P. Finkbeiner, and G. Dobler, *Phys. Rev.* **D76**, 083012 (2007), 0705.3655.
 - [8] G. Dobler and D. P. Finkbeiner, *Astrophys. J.* **680**, 1222 (2008), 0712.1038.
 - [9] G. Dobler, D. P. Finkbeiner, I. Cholis, T. R. Slatyer, and N. Weiner (2009), 0910.4583.
 - [10] M. Cirelli, M. Kadastik, M. Raidal, and A. Strumia, *Nucl. Phys.* **B813**, 1 (2009), 0809.2409.
 - [11] D. P. Finkbeiner and N. Weiner, *Phys. Rev.* **D76**, 083519 (2007), astro-ph/0702587.
 - [12] N. Arkani-Hamed, D. P. Finkbeiner, T. R. Slatyer, and N. Weiner, *Phys. Rev.* **D79**, 015014 (2009), 0810.0713.
 - [13] M. Pospelov and A. Ritz, *Phys. Lett.* **B671**, 391 (2009), 0810.1502.
 - [14] I. Cholis, D. P. Finkbeiner, L. Goodenough, and N. Weiner (2008), 0810.5344.
 - [15] I. Cholis, G. Dobler, D. P. Finkbeiner, L. Goodenough, and N. Weiner (2008), 0811.3641.
 - [16] J. T. Ruderman and T. Volansky, *JHEP* **02**, 024 (2010), 0908.1570.
 - [17] K. R. Dienes, C. F. Kolda, and J. March-Russell, *Nucl. Phys.* **B492**, 104 (1997), hep-ph/9610479.
 - [18] C. Cheung, J. T. Ruderman, L.-T. Wang, and I. Yavin (2009), 0902.3246.
 - [19] A. Katz and R. Sundrum (2009), 0902.3271.
 - [20] D. E. Morrissey, D. Poland, and K. M. Zurek (2009), 0904.2567.
 - [21] Y. Cui, D. E. Morrissey, D. Poland, and L. Randall, *JHEP* **05**, 076 (2009), 0901.0557.
 - [22] R. Bernabei et al. (DAMA), *Phys. Lett.* **B480**, 23 (2000).
 - [23] R. Bernabei et al., *Int. J. Mod. Phys.* **D13**, 2127 (2004), astro-ph/0501412.
 - [24] R. Bernabei et al. (DAMA), *Eur. Phys. J.* **C56**, 333 (2008), 0804.2741.
 - [25] R. Bernabei et al. (2010), 1002.1028.

- [26] C. E. Aalseth et al. (CoGeNT) (2010), 1002.4703.
- [27] A. L. Fitzpatrick, D. Hooper, and K. M. Zurek (2010), 1003.0014.
- [28] S. Andreas, C. Arina, T. Hambye, F.-S. Ling, and M. H. G. Tytgat (2010), 1003.2595.
- [29] P. Graham, R. Harnik, S. Rajendran, and P. Saraswat (hep-ph preprint April 5, 2010: “Exothermic Dark Matter”).
- [30] A. W. Strong et al., *Astron. Astrophys.* **444**, 495 (2005), astro-ph/0509290.
- [31] D. E. Kaplan, M. A. Luty, and K. M. Zurek, *Phys. Rev.* **D79**, 115016 (2009), 0901.4117.
- [32] E. Kuflik, A. Pierce, and K. M. Zurek (2010), 1003.0682.
- [33] J. Giedt, A. W. Thomas, and R. D. Young, *Phys. Rev. Lett.* **103**, 201802 (2009), 0907.4177.
- [34] R. Essig, P. Schuster, and N. Toro (2009), 0903.3941.
- [35] B. Batell, M. Pospelov, and A. Ritz (2009), 0903.0363.
- [36] F. Bossi (2009), 0904.3815.
- [37] G. Amelino-Camelia et al. (2010), 1003.3868.
- [38] M. Freytsis, Z. Ligeti, and J. Thaler (2009), 0911.5355.
- [39] K. L. McDonald and D. E. Morrissey (2010), 1002.3361.
- [40] B. Aubert (The BABAR) (2009), 0902.2176.
- [41] P.-f. Yin, J. Liu, and S.-h. Zhu (2009), 0904.4644.
- [42] H.-B. Li and T. Luo, *Phys. Lett.* **B686**, 249 (2010), 0911.2067.
- [43] S. Chang, A. Pierce, and N. Weiner, *Phys. Rev.* **D79**, 115011 (2009), 0808.0196.
- [44] J. Angle et al. (XENON10), *Phys. Rev.* **D80**, 115005 (2009), 0910.3698.
- [45] A. Manzur et al. (2009), 0909.1063.
- [46] J. Angle et al. (XENON), *Phys. Rev. Lett.* **100**, 021303 (2008), 0706.0039.
- [47] J. Filippini, Talk at *Les Recontres de Physique de la Vallée d’Aosta* <http://agenda.infn.it/getFile.py/access?contribId=10&sessionId=2&resId=0&materialId=slides&confId=930>.
- [48] E. M. Drobyshevski, *Mod. Phys. Lett.* **A23**, 3077 (2008), 0706.3095.
- [49] R. Bernabei et al., *Eur. Phys. J.* **C53**, 205 (2008), 0710.0288.
- [50] J. March-Russell, C. McCabe, and M. McCullough, *JHEP* **05**, 071 (2009), 0812.1931.
- [51] D. P. Finkbeiner, T. R. Slatyer, N. Weiner, and I. Yavin, *JCAP* **0909**, 037 (2009), 0903.1037.
- [52] D. S. M. Alves, S. R. Behbahani, P. Schuster, and J. G. Wacker (2009), 0903.3945.
- [53] D. S. M. Alves, S. R. Behbahani, P. Schuster, and J. G. Wacker (2010), 1003.4729.

- [54] K. M. Zurek, Phys. Rev. **D79**, 115002 (2009), 0811.4429.
- [55] K. Jedamzik, Phys. Rev. **D74**, 103509 (2006), hep-ph/0604251.
- [56] A. R. Pullen, R.-R. Chary, and M. Kamionkowski, Phys.Rev. **D76**, 063006 (2007), astro-ph/0610295.
- [57] L. Randall and R. Sundrum, Nucl. Phys. **B557**, 79 (1999), hep-th/9810155.
- [58] G. F. Giudice, M. A. Luty, H. Murayama, and R. Rattazzi, JHEP **12**, 027 (1998), hep-ph/9810442.
- [59] I. Cholis and N. Weiner (2009), 0911.4954.
- [60] J. L. Feng, M. Kaplinghat, and H.-B. Yu (2009), 0911.0422.
- [61] P. Meade, S. Nussinov, M. Papucci, and T. Volansky (2009), 0910.4160.
- [62] R. Essig, N. Sehgal, and L. E. Strigari (2009), 0902.4750.
- [63] A. A. Abdo, M. Ackermann, and M. Ajello (2010), 1001.4836.
- [64] M. Kawasaki, K. Kohri, and T. Moroi, Phys. Rev. **D71**, 083502 (2005), astro-ph/0408426.
- [65] D. Eichler, Phys. Rev. Lett. **63**, 2440 (1989).
- [66] E. Nardi, F. Sannino, and A. Strumia, JCAP **0901**, 043 (2009), 0811.4153.
- [67] A. Arvanitaki et al. (2008), 0812.2075.
- [68] R. Bernabei et al. (DAMA), Mod. Phys. Lett. **A23**, 2125 (2008), 0802.4336.
- [69] C. E. Aalseth et al. (CoGeNT), Phys. Rev. Lett. **101**, 251301 (2008), 0807.0879.
- [70] E. M. Drobyshevski (2007), arXiv/0706.3095.
- [71] M. Fairbairn and T. Schwetz, JCAP **0901**, 037 (2009), 0808.0704.
- [72] G. B. Gelmini, J. Phys. Conf. Ser. **203**, 012042 (2010), 0910.3032.
- [73] G. B. Gelmini, P. Gondolo, and N. Bozorgnia.
- [74] R. Bernabei et al. (DAMA), Nucl. Instrum. Meth. **A592**, 297 (2008), 0804.2738.
- [75] Z. Ahmed et al. (The CDMS-II) (2009), 0912.3592.
- [76] D. S. Akerib et al. (CDMS), Phys. Rev. Lett. **96**, 011302 (2006), astro-ph/0509259.
- [77] B. Aubert et al. (BABAR) (2009), 0908.2821.
- [78] B. Aubert et al. (BABAR), Phys. Rev. Lett. **103**, 251801 (2009), 0908.2840.
- [79] T. C. Collaboration (2010), 1003.0417.
- [80] B. Batell, M. Pospelov, and A. Ritz (2009), 0911.4938.
- [81] M. J. Strassler and K. M. Zurek, Phys. Lett. **B651**, 374 (2007), hep-ph/0604261.
- [82] M. Baumgart, C. Cheung, J. T. Ruderman, L.-T. Wang, and I. Yavin, JHEP **04**, 014 (2009),

0901.0283.

- [83] C. Cheung, J. T. Ruderman, L.-T. Wang, and I. Yavin (2009), 0909.0290.
- [84] V. M. Abazov et al. (D0), Phys. Rev. Lett. **103**, 081802 (2009), 0905.1478.
- [85] A. Falkowski, J. T. Ruderman, T. Volansky, and J. Zupan (2010), 1002.2952.
- [86] J. D. Bjorken, R. Essig, P. Schuster, and N. Toro, Phys. Rev. **D80**, 075018 (2009), 0906.0580.
- [87] M. Reece and L.-T. Wang (2009), 0904.1743.
- [88] B. Batell, M. Pospelov, and A. Ritz, Phys. Rev. **D80**, 095024 (2009), 0906.5614.
- [89] R. Essig, P. Schuster, N. Toro, and B. Wojtsekhowski (2010), 1001.2557.
- [90] M. Freytsis, G. Ovanesyan, and J. Thaler (2009), 0909.2862.
- [91] S. Heinemeyer, Y. Kahn, M. Schmitt, and M. Velasco (2007), 0705.4056.
- [92] P. Schuster, N. Toro, and I. Yavin, Phys. Rev. **D81**, 016002 (2010), 0910.1602.
- [93] E. Abouzaid et al. (KTeV), Phys. Rev. **D75**, 012004 (2007), hep-ex/0610072.
- [94] B. Batell, M. Pospelov, A. Ritz, and Y. Shang (2009), 0910.1567.
- [95] P. Schuster, N. Toro, N. Weiner, and I. Yavin (2009), 0910.1839.
- [96] N. Arkani-Hamed and N. Weiner, JHEP **12**, 104 (2008), 0810.0714.
- [97] S. Chang, J. Liu, A. Pierce, N. Weiner, and I. Yavin (2010), ??
- [98] J. D. Lewin and P. F. Smith, Astropart. Phys. **6**, 87 (1996).
- [99] M. J. Reid, K. M. Menten, A. Brunthaler, and G. A. Moellenbrock (2009), 0902.3928.
- [100] M. C. Smith et al., Mon. Not. Roy. Astron. Soc. **379**, 755 (2007), astro-ph/0611671.
- [101] W. Dehnen and J. Binney, Mon. Not. Roy. Astron. Soc. **298**, 387 (1998), astro-ph/9710077.
- [102] D. Tucker-Smith and N. Weiner, Phys. Rev. **D64**, 043502 (2001), hep-ph/0101138.