

# Testing the Friedmannian Magnitude-redshift Relation with SNIa Data

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Standard cosmology is constructed upon the (generally implicit) assumption of the “large scale” homogeneity of our Universe. Now, structures are observed at scales which become larger and larger as the observational distances increase. However, the homogeneous Friedmann-Lemaître-Robertson-Walker model remains a cosmological paradigm, and Friedmannian relations are usually used to work out the model, implicitly assuming that the homogeneity assumption is valid at the studied scales. This has been the way the concordance ( $\Lambda$ CDM) model came out from the analyses of the SNIa data. Since this model implies the validity of the Friedmannian magnitude-redshift relation at the range of redshifts spanned by the supernovae surveys, it has been proposed some years ago a very simple test of this possible validity, provided the SNIa should be confirmed as standard candles. Since we can hope that the redshifts of the supernovae which will be observed in the years to come will allow such a test to be performed, it would be interesting to find inside the SNIa community some collaborators ready to perform the test.

## 1. The three observational pillars of the standard Hot Big-Bang model

The standard Hot Big-Bang model is based upon three observational pillars, of which the most common interpretations have been thoroughly discussed in a review paper published in 2003 by López-Corredoira [1]. We refer here to this author, even if some of the possible caveats proposed in his work seem too weak to really shake the actual robustness of these pillars, as far as our current scientific knowledge is involved.

1. The increasing with distance galaxy redshifts are commonly interpreted as due to a recession of the galaxies implying an expanding Universe. This interpretation seems to be the robustest available upon the market even if some authors have proposed other ones.
2. The measured abundance of the light elements is generally explained in the framework of the standard Big Bang primordial nucleosynthesis. Alternatives can be found in the literature, although this model might be considered as the best proposed up to now, besides some possible problems that might be or not solved in the future.
3. The cosmic microwave background (CMB) is mostly believed to be a relic of the early stage of the Universe, when the photon decoupled from the high energy primordial soup, that is observed as a background with a perfect black-body spectrum. By now, nobody has proposed a satisfactory alternative scenario and therefore the standard interpretation remains the best one. However, López-Corredoira [1] points out, for example, a possible problem arising from the high energy cosmic rays. These are currently measured with energies up to  $3.2 \cdot 10^{20}$  eV, which is, in the framework of standard physics

(another explanation has been proposed in the framework of Scale Relativity [2]) beyond the theoretical energy limit for particles traveling for distances such as those separating powerful galaxies, because of their interactions with the CMB. Considering the robustness of the other observations in favor of the standard interpretation of the CMB in the Hot Big-Bang model, it seems more clever to consider that this problem pertains to particle physics rather than to cosmology.

## 2. The “concordance” $\Lambda$ CDM model

The current program of a major part of the cosmological community is to determine a set of cosmological “constants” in the framework of a Friedmann-Lemaître-Robertson-Walker (FLRW) model of Universe incorporating an inflationary scenario.

- The Hubble Space Telescope (HST) provides observations allowing an estimate of the Hubble constant  $H_0$ .
- The deuterium abundance measured in Lyman  $\alpha$  clouds gives a value for  $\Omega_b$ .
- The abundance of rich clusters in large scale structures leads to an estimate for  $\sigma_8$ .
- Large scale weak lensing is a probe for  $\Omega_M$  and  $w$ .
- The magnitude-redshift relation issuing from the SNIa observations provides ranges for  $H_0$ ,  $\Omega_M$  and  $\Omega_\Lambda$ .
- The CMB  $C_l$  curve is used to put constraints on 6 to 12 cosmological “constants”, the number of constants depending on the cosmological assumptions retained.

A cross-correlation of all these results, and others that cannot be cited here, an exhaustive review of the cosmological field being beyond the scope of this contribution, provides what is known as the concordance or  $\Lambda$ CDM model, with the following values of the main cosmological “constants” pertaining to this model (see, e.g., [3]) :

$H_0 \sim 70 \text{ km/s/Mpc}$ ,  $\Omega_b \sim 0.02h^{-2}$ ,  $\Omega_0 \sim 1$ ,  $\Omega_M \sim 0.3$ ,  $\Omega_\Lambda \sim 0.7$ ,  $n \sim 1$ ,  $\tau \sim 0$ , etc.

## 2.1. Priors to the “concordance” model:

However, the “concordance” model is built with priors pertaining to the “standard” paradigm of cosmology and which are:

- The FLRW metric of the Universe.  
Since all isotropic, homogeneous, uniformly expanding universes exhibit a FLRW metric [4], their geometry can be characterized by a finite number of cosmological “constants”.  
It is worth stressing here that the FLRW solution implies, as an implicit assumption, that global solutions to Einstein’s equations can be physical, although we know that General Relativity is an essentially local theory. A lot of work can be found in the literature (for a review see, e.g., [6]) which discuss the validity of such an approach. In the FLRW case, the stress-energy tensor is averaged over a scale which is generally neither precised nor probed. However, there is currently no available mathematical definition of a tensor average. Averaging is only mathematically valid for scalars. Therefore, Friedmannian relations must be used with great care and should be considered as approximations and systematically validated for the scale range to which they are applied.
- The inflation assumption.  
It implies a power-law primordial spectrum of the density perturbations presumably generated during an inflationary phase and assumed for the analysis of the CMB temperature fluctuations. Another rather generic prediction of this paradigm is a strict spatial flatness of the Universe, implying  $\Omega_0 = 1$ . The current error bars on  $\Omega_0$  ( $\Omega_0 = 1.02 \pm 0.02$  [5]) are still too large to test this prediction, which might be more easy to discard than to validate, since the prediction is  $\Omega_0 =$  (strictly) 1.

## 2.2. Other theoretical alternatives

Other theoretical alternatives have been proposed by different authors. The reader will find below some among the most famous examples:

- Gaussian adiabatic fluctuations versus isocurvature
- $\Lambda$ CDM with a cosmological constant versus quintessence
- vanishing versus non vanishing amount of gravitational waves
- vanishing versus non vanishing cosmic string component
- etc.

## 3. The homogeneity assumption

“Large scale spatial homogeneity” is generally assumed without any precision as to the scale at which the Universe is supposed to become homogeneous. However, large scale inhomogeneities can yield a geometry noticeably different from the one inferred from the homogeneous assumption, therefore invalidating the use of Friedmannian expressions. Inhomogeneous solutions to Einstein’s equations generally involve cosmological functions of the coordinates rather than cosmological constants as in FLRW solutions (see, e.g., for a mere example, [7]).

As a consequence, the expressions or relations retained to analyse data collected at any distances must be carefully used, in particular those assuming the “constance” of the cosmological “constants”.

But these expressions can be used to test some of the features of the geometry of the Universe. This is, in particular, the case of the FLRW expression for the luminosity distance as a function of redshift and of the three “constants”  $H_0$ ,  $\Omega_M$  and  $\Omega_\Lambda$ , which has been initially retained to analyse the data of the SNIa [8, 9].

## 4. Use of the magnitude redshift-relation to test some features of the geometry of the Universe

The FLRW model implies the use of a relation for the luminosity distance  $D_L$  as a function of the four parameters,  $z$ ,  $H_0$ ,  $\Omega_M$  and  $\Omega_\Lambda$ . This is only valid if the FLRW picture is confirmed as a good approximation for all the spanned values of  $z$ . The relation is the following:

$$D_L = \frac{c(1+z)}{H_0 \sqrt{|\kappa|}} \mathcal{S} \left( \sqrt{|\kappa|} \int_0^z [(1+z')^2(1+\Omega_M z') - z'(2+z')\Omega_\Lambda]^{-\frac{1}{2}} dz' \right), \quad (1)$$

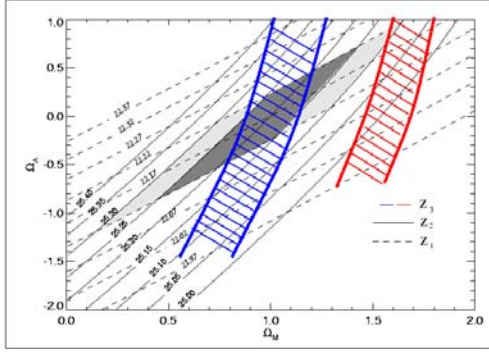


Figure 1: Contours of constant  $m$  in the  $\Omega_M$ - $\Omega_\Lambda$  plane.

for  $\Omega_M + \Omega_\Lambda > 1$   $\mathcal{S} = \sin$  and  $\kappa = 1 - \Omega_M - \Omega_\Lambda$   
 for  $\Omega_M + \Omega_\Lambda < 1$   $\mathcal{S} = \sinh$  and  $\kappa = 1 - \Omega_M - \Omega_\Lambda$   
 for  $\Omega_M + \Omega_\Lambda = 1$   $\mathcal{S} = I$  and  $\kappa = 1$ .

The apparent magnitude of a source at a given  $z$  writes

$$m = M + 5 \log D_L(z; H_0, \Omega_M, \Omega_\Lambda) + 25. \quad (2)$$

The proposed method, inspired from Goobar and Perlmutter, 1995 [10] uses this expression to draw contours of constant apparent magnitude on the  $\Omega_\Lambda$ -versus- $\Omega_M$  plane for different fixed redshifts. One first selects, from the SNIa catalogs, three fair samples at sufficiently different redshifts,  $z_1 < z_2 < z_3$ . Sets of rather low redshifts can be used to try to find out the scale of a possible transition to homogeneity, usually expected above some hundred Mpc. For the higher reached  $z_3$ , this transition, if not found at lower redshifts, may also be put to the test. Note that the errors allowed for small redshift samples used for this kind of analysis would be lower than for high redshift ones.

The uncertainties in the measurements define three strips between two contour lines for each allowed range of  $\Omega_M$  and  $\Omega_\Lambda$ . The two first ones ( $z \sim z_1$ ,  $z \sim z_2$ ) cross, thus defining a possible allowed region (dashed rhombus in 1). The third strip ( $z \sim z_3$ ) does or does not cross this rhombus.

#### 4.1. Possible missing of the rhombus

If the data produce the red case in Figure 1, where the third strip does not cross the rhombus, the FLRW magnitude-redshift relation might be discarded, at least at the spanned redshifts.

As a consequence, one (or more) of the assumptions retained to derive this relation must be removed: either homogeneity, or the constance of the Cosmological Constant, or the universe expansion (but this last seems the robuster one, being one of the three pillars of the standard model of universe, see above

Section 1) etc.

#### 4.2. Possible crossing of the rhombus

If the data yield the blue strip case of Figure 1, this would imply a validation of the FLRW magnitude-redshift relation at the spanned redshifts.

As a consequence, this would be rather good support for the homogeneity assumption, for the use of this method to obtain the value of  $\Omega_\Lambda$  and for its interpretation as a real Cosmological Constant versus quintessence.

Note that the “concordance” value  $\Omega_\Lambda \sim 0.7$  corresponds to a theoretical prediction of the value of the Cosmological Constant in the framework of the scale relativity theory [2, 11].

But one should remain careful since it has been shown that inhomogeneous cosmological models, with or without a non-zero cosmological constant, can mimic FLRW ones with a different value for  $\Lambda$  [7, 12].

#### 5. Conclusion

- The current “concordance” model of the Universe is a very interesting product of crossed analyses of the large amount of currently available cosmological data.
- We must however remain aware of its priors, mostly implicitly assumed to derive the geometry and the dynamics of our Universe with those methods.
- The Friedmannian magnitude-redshift relation can be used as a test, using fair samples of SNIa data at different redshifts to: i) determine the redshift value of the possible transition from inhomogeneity to homogeneity; ii) test the Cosmological Constant versus quintessence or other hypotheses; iii) and more generally the FLRW model at the spanned redshifts.
- It could be also an independent way of testing the validity of the analysis of the fractal dimension of the galaxy distribution currently in process [15], by comparing the values for this transition obtained by both methods.

One must however remain aware of the high degeneracy of the cosmological exercise. When dealing with only one Universe, many different answers can be proposed to a given problem. The application of Occam’s razor to discriminate between a number of

abundant ideas, some of them very ingenious and appealing, is the usual and more economic way to obtain at least an elegant, at most “the” right solution. This is the reason why the proposed test retains a set of very simple assumptions in the framework of Einstein’s General Relativity. Of course, other interpretations than those proposed in Section 4 can be put forward to explain the results of the test. The high degeneracy of the cosmological problem, here stressed once again, allows each reader to choose among the possible explanations the one which corresponds to his (her) own taste, provided not too many epicycles would be needed to construct his (her) model of Universe (do not forget Occam’s razor).

## 6. Prospects

To complete the proposed test we need:

- A validation of the SNIa as accurate standard candles, actively in progress inside the SNIa community [13, 14].
- The largest catalogs of photo-spectroscopic measurements of SNIa from very low (needed for calibration) to sufficiently high redshifts that can or will be found in the on-going or planned surveys: the Nearby Supernovae Factory (SNfactory), the Supernova Legacy Survey (SNLS), Equation of State: SuperNovae trace Cosmic Expansion (ESSENCE), the Supernova/Acceleration Probe (SNAP), etc.
- A collaboration with the SNIa teams to use the more recent data to perform a statistically robust analysis.

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